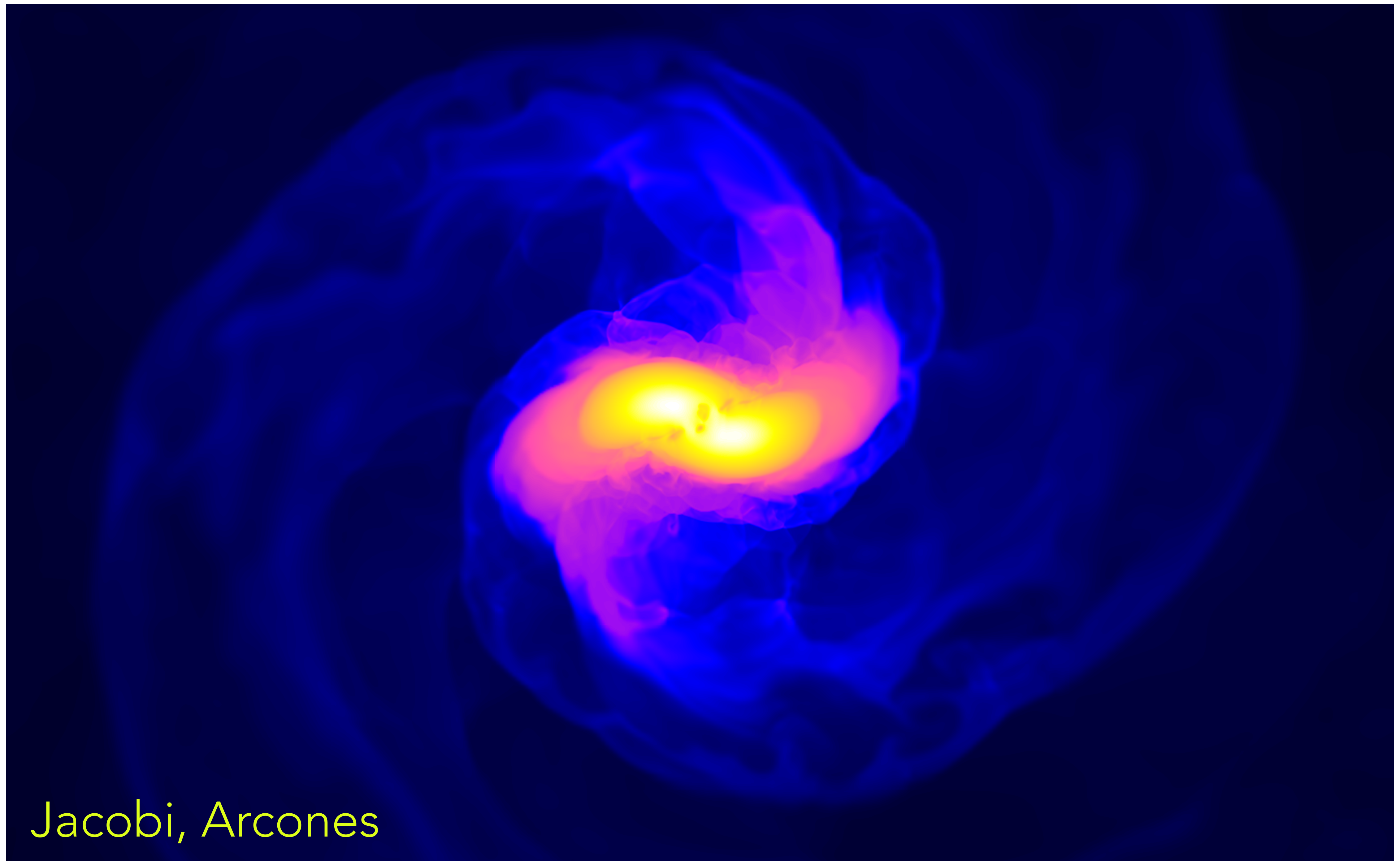
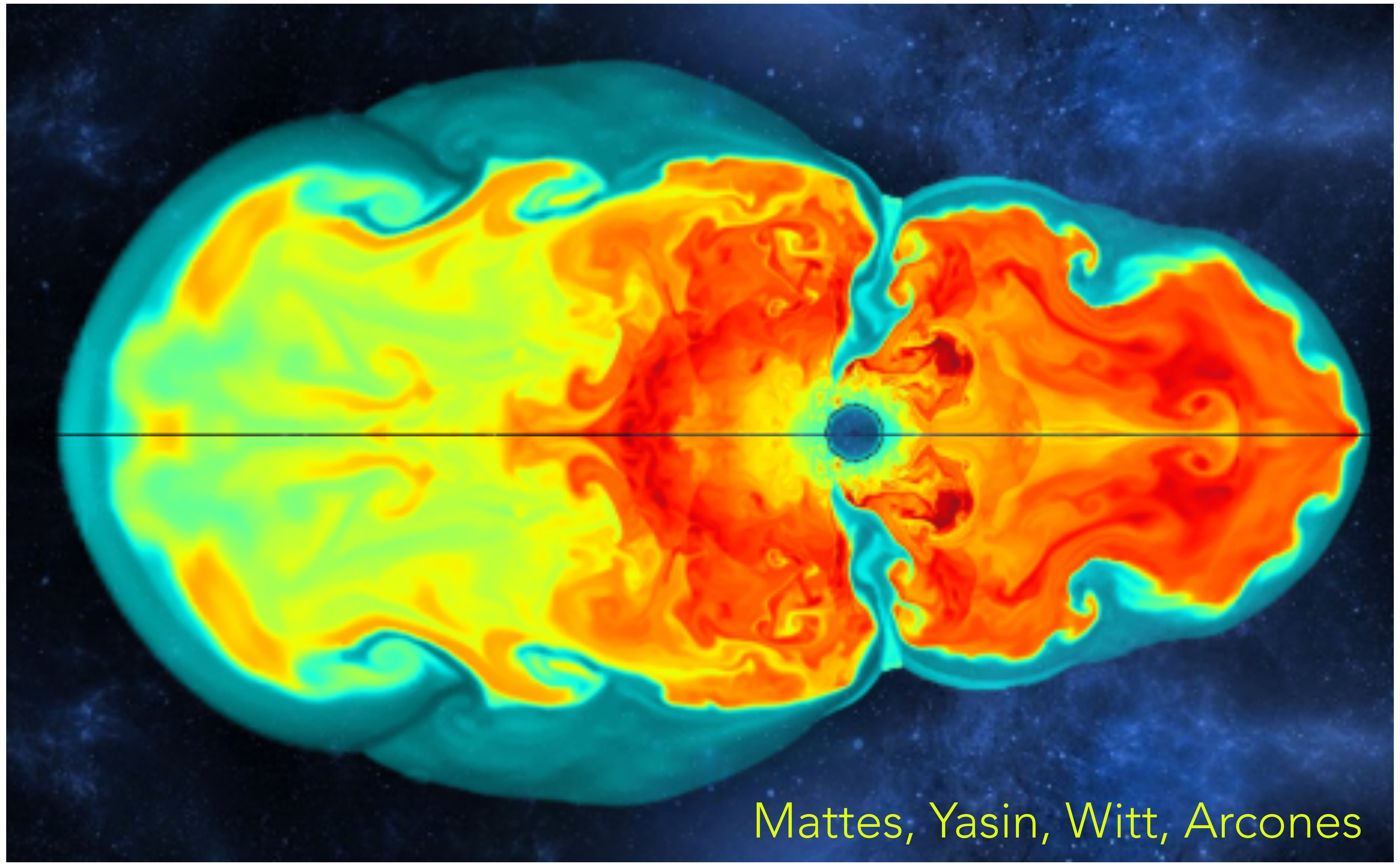
Nuclear Physics in Astrophysics XI
Dresden, 15-20 Sept. 2024

R-process in neutron star mergers and supernovae



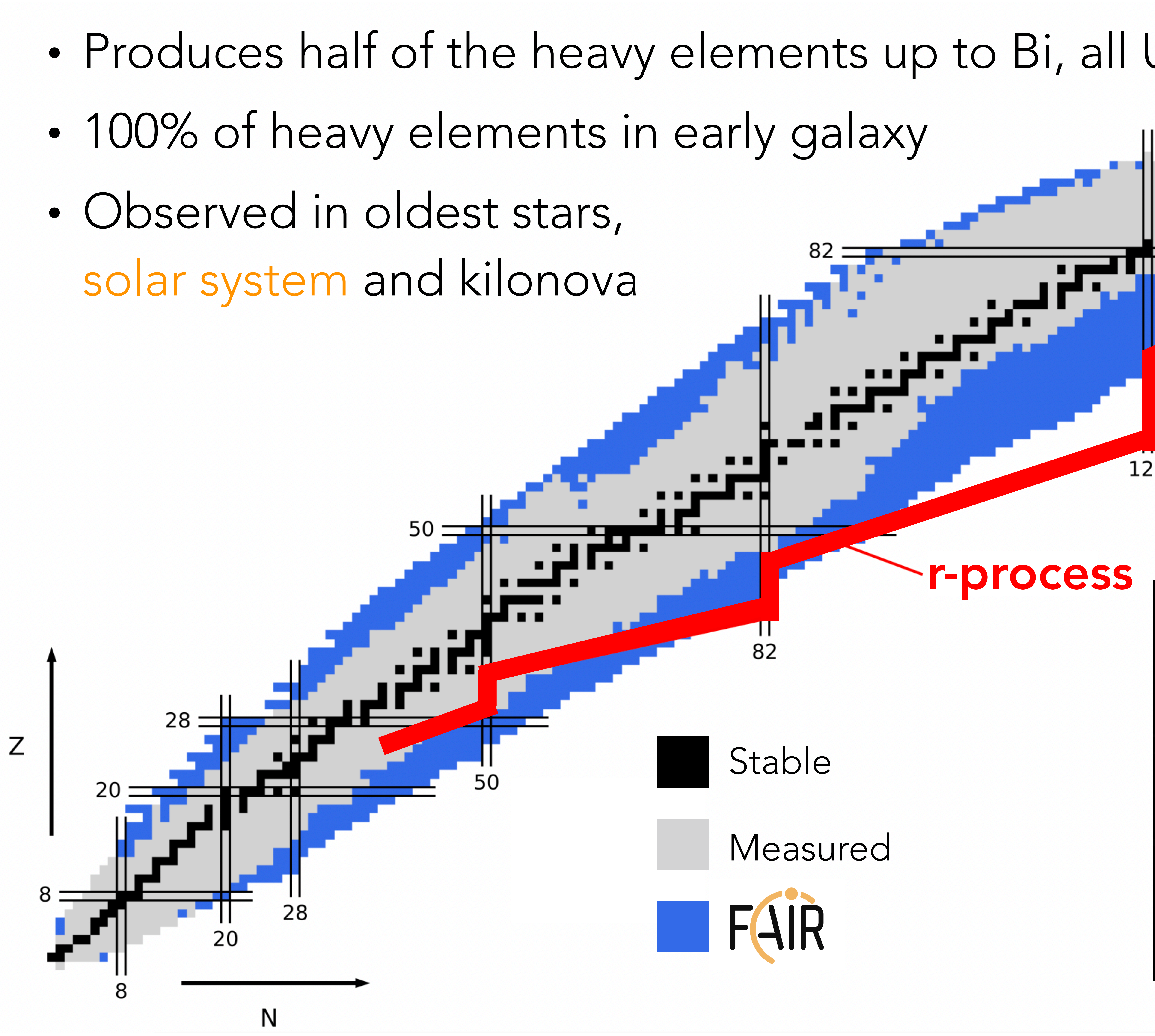


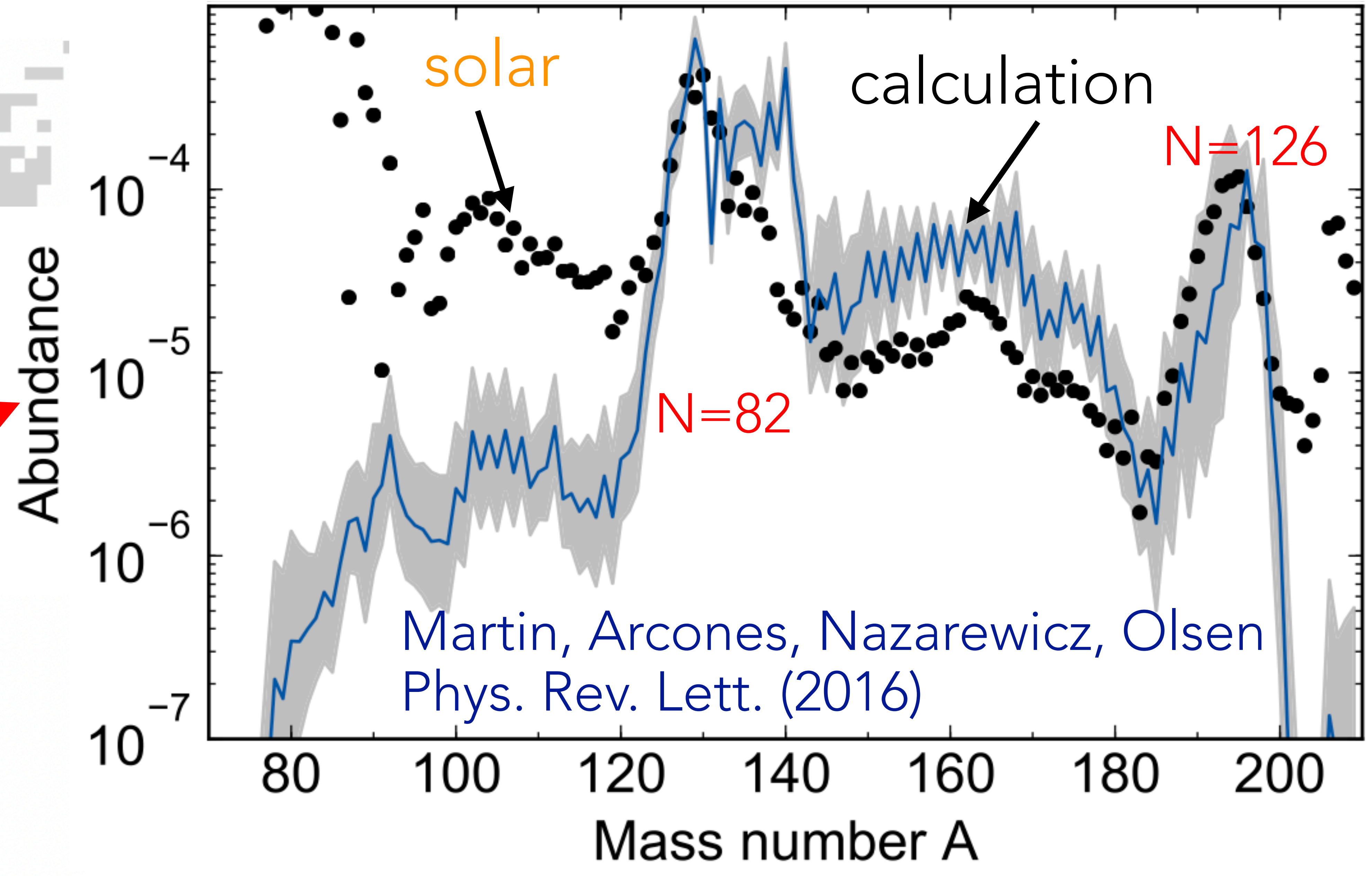
Almudena Arcones

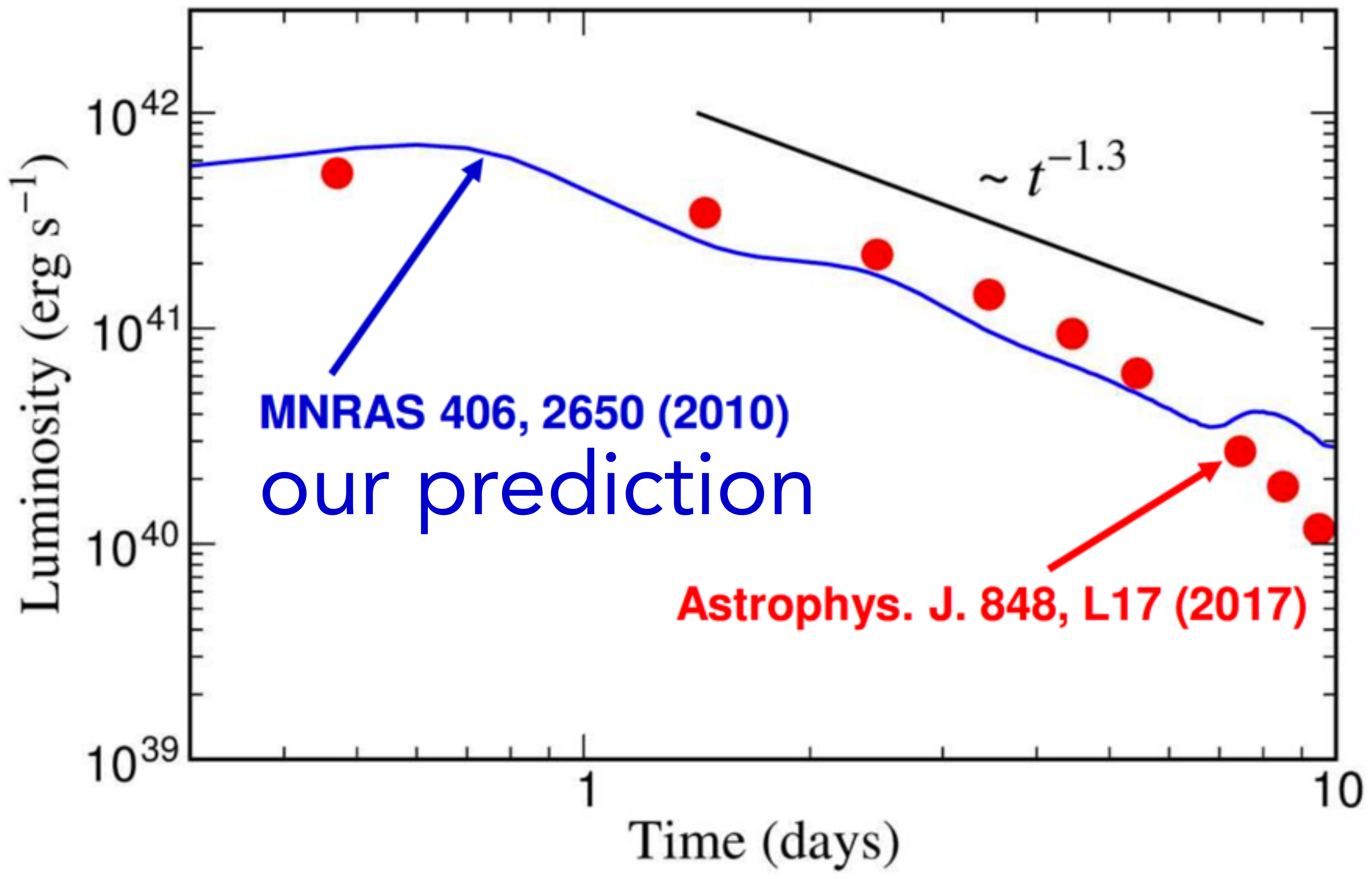
Max-Planck-Institut für Kernphysik Heidelberg
Technische Universität Darmstadt
GSI

Rapid neutron capture process

- Produces half of the heavy elements up to Bi, all U and Th
- 100% of heavy elements in early galaxy
- Observed in oldest stars, solar system and kilonova





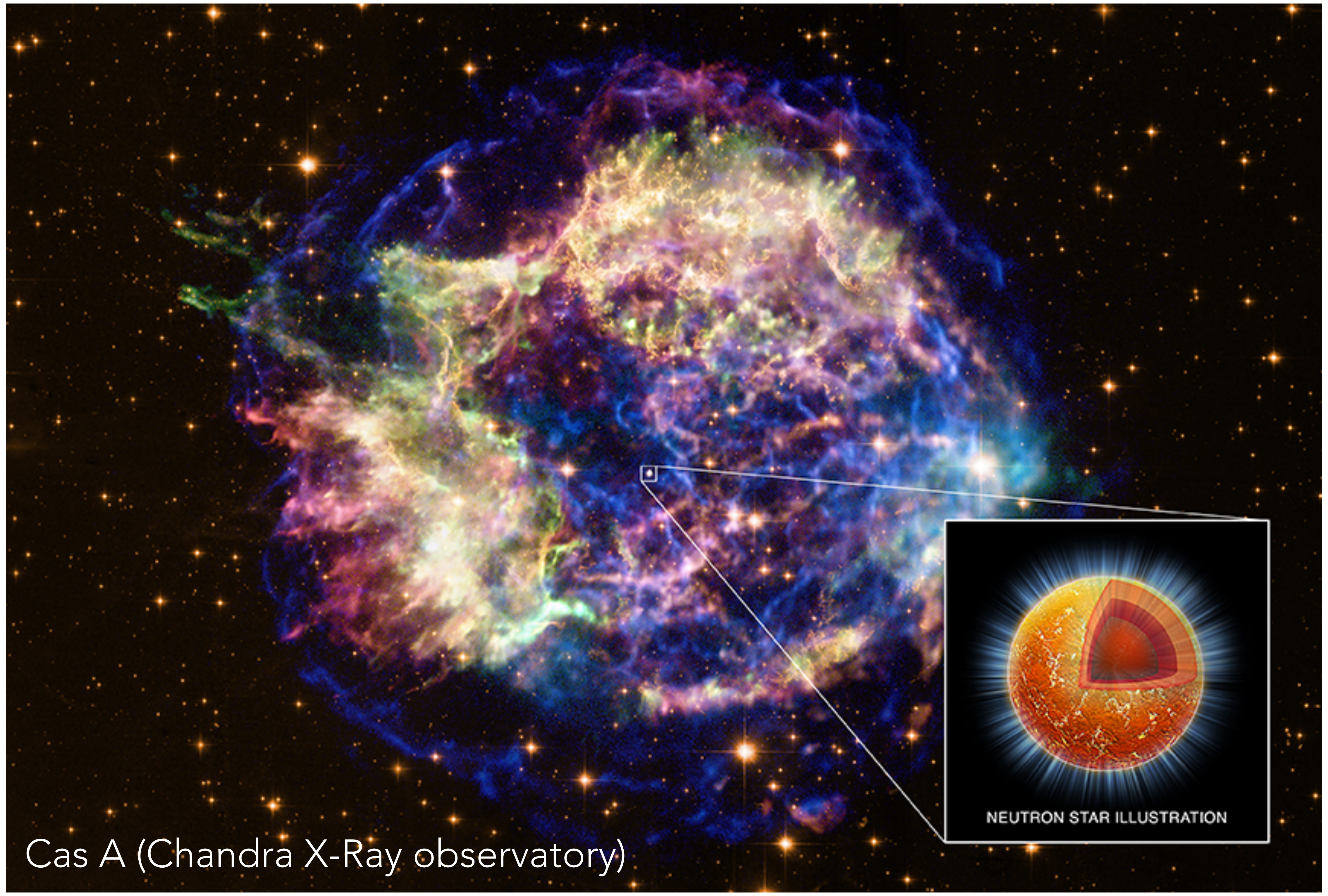


Origin of heavy elements?

Rapid neutron capture process

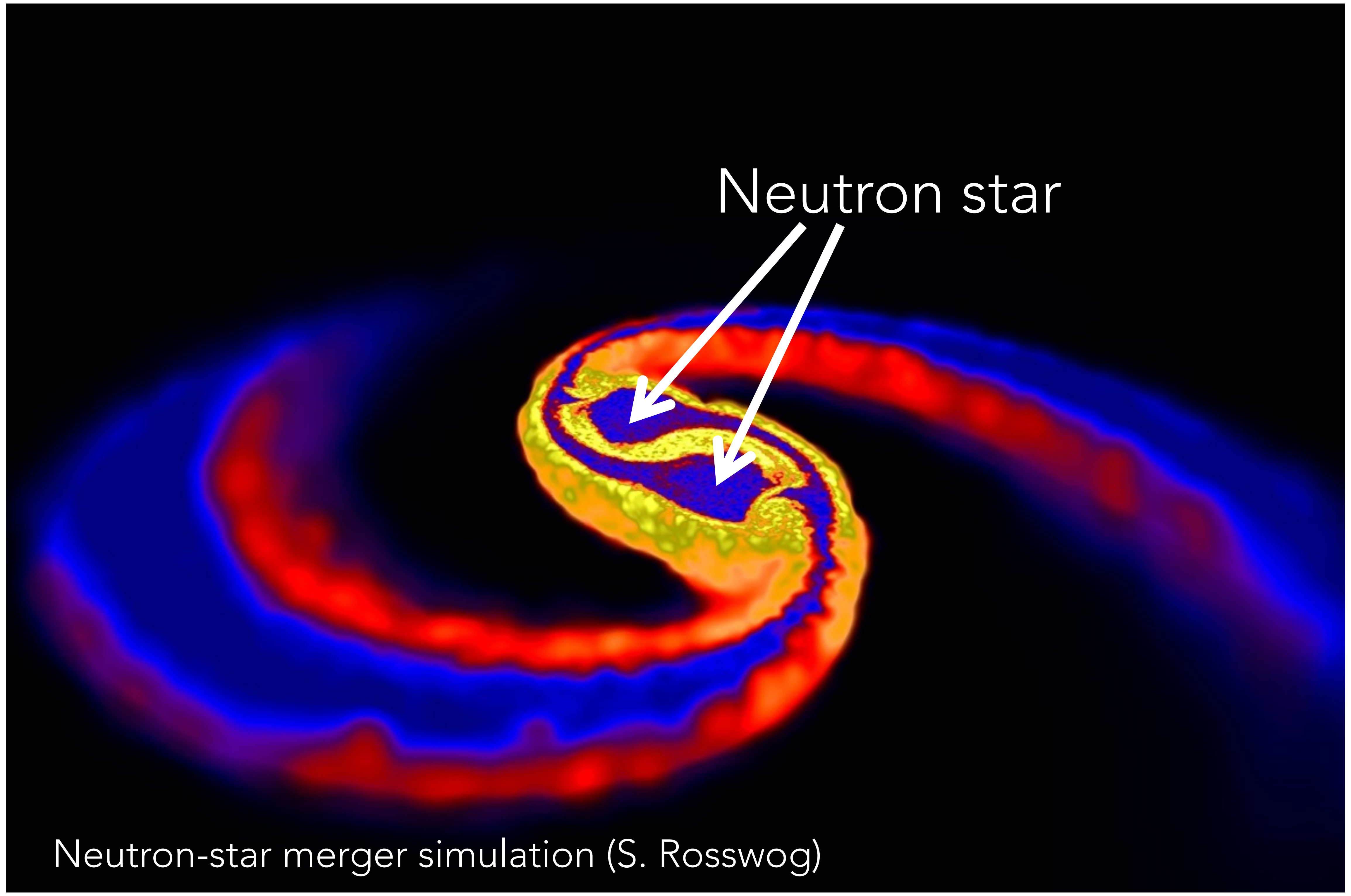
Explosive and high neutron densities

Rare Supernovae



Cas A (Chandra X-Ray observatory)

Neutron star mergers

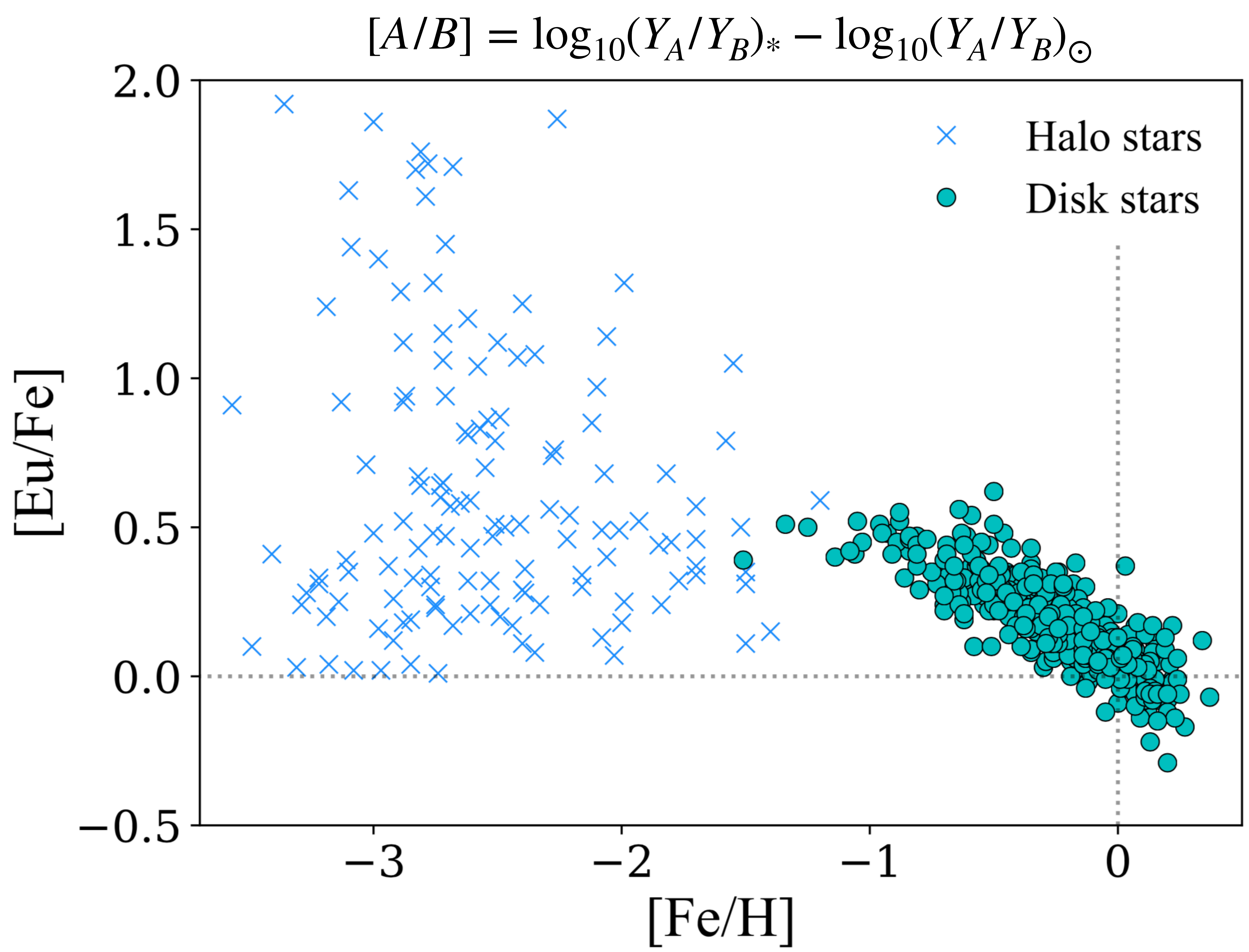


Neutron-star merger simulation (S. Rosswog)

Observations and galactic chemical evolution

Evolution with time (or metallicity) -> Galactic Chemical Evolution (GCE)

-> r-process sites: mergers vs. supernovae



Matteucci et al. MNRAS (2014), Côté et al. ApJ (2019), Molero et al. MNRAS (2021)

Observations and galactic chemical evolution

Evolution with time (or metallicity) -> Galactic Chemical Evolution (GCE)

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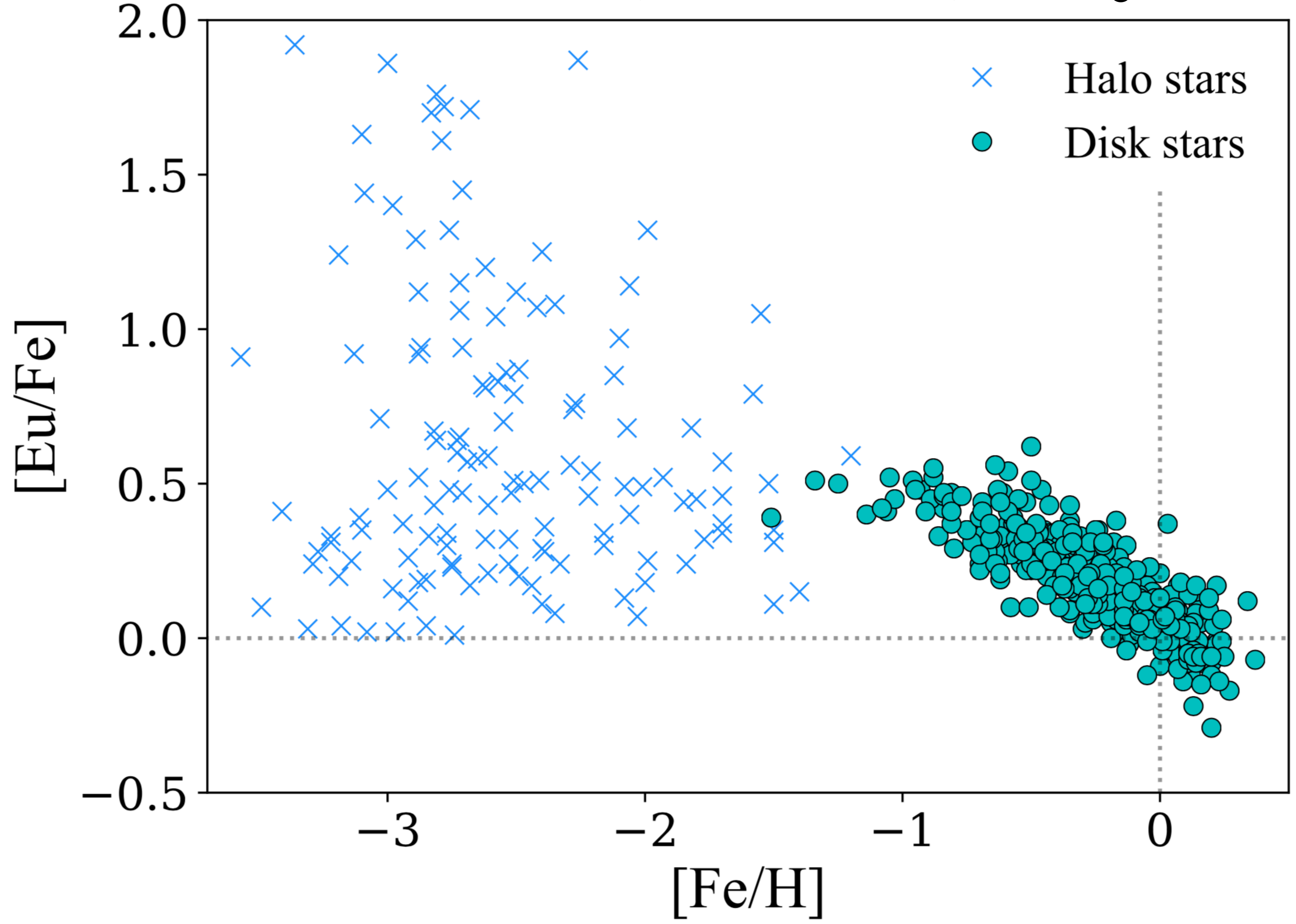
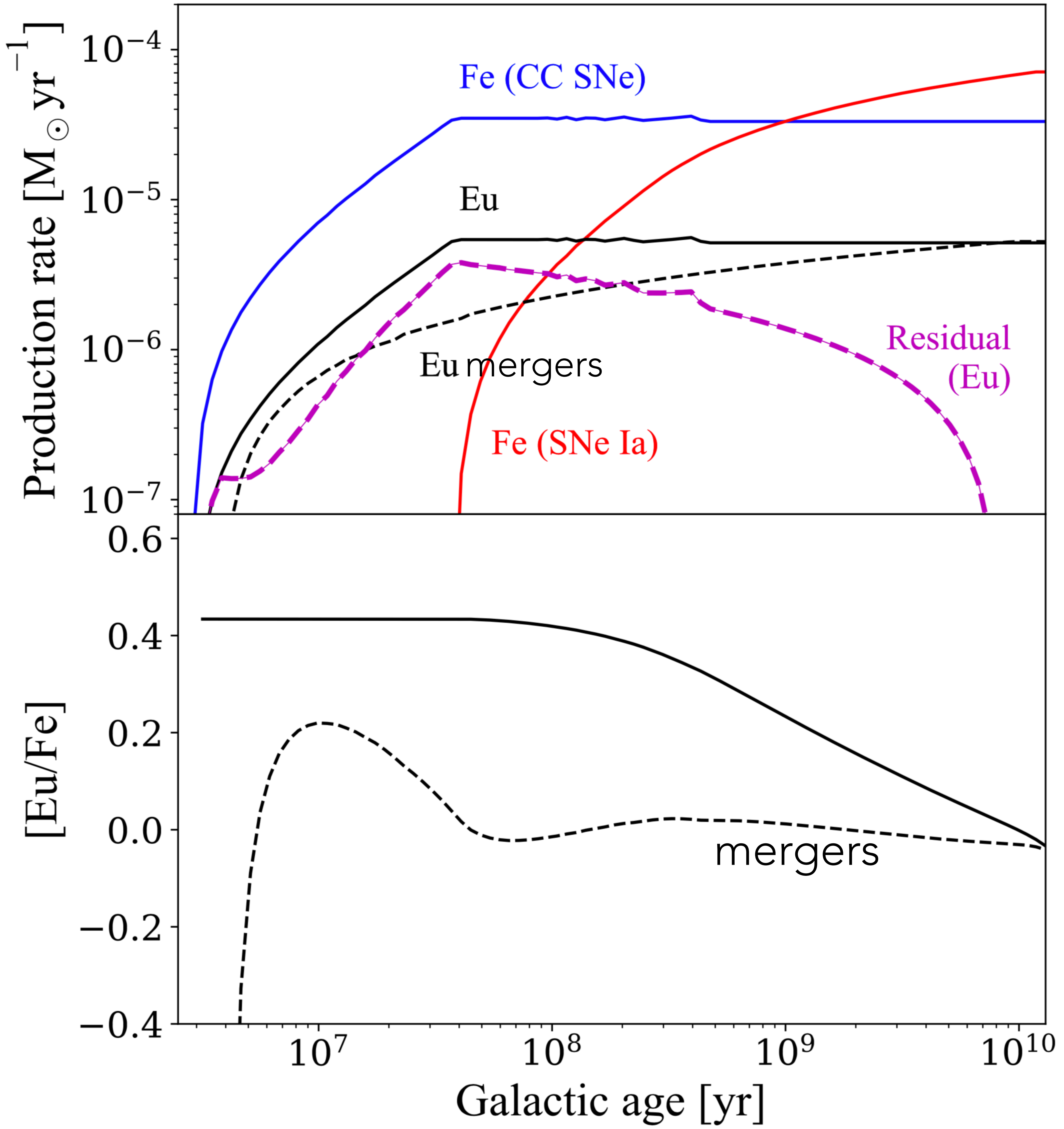
$$[A/B] = \log_{10}(Y_A/Y_B)_* - \log_{10}(Y_A/Y_B)_\odot$$

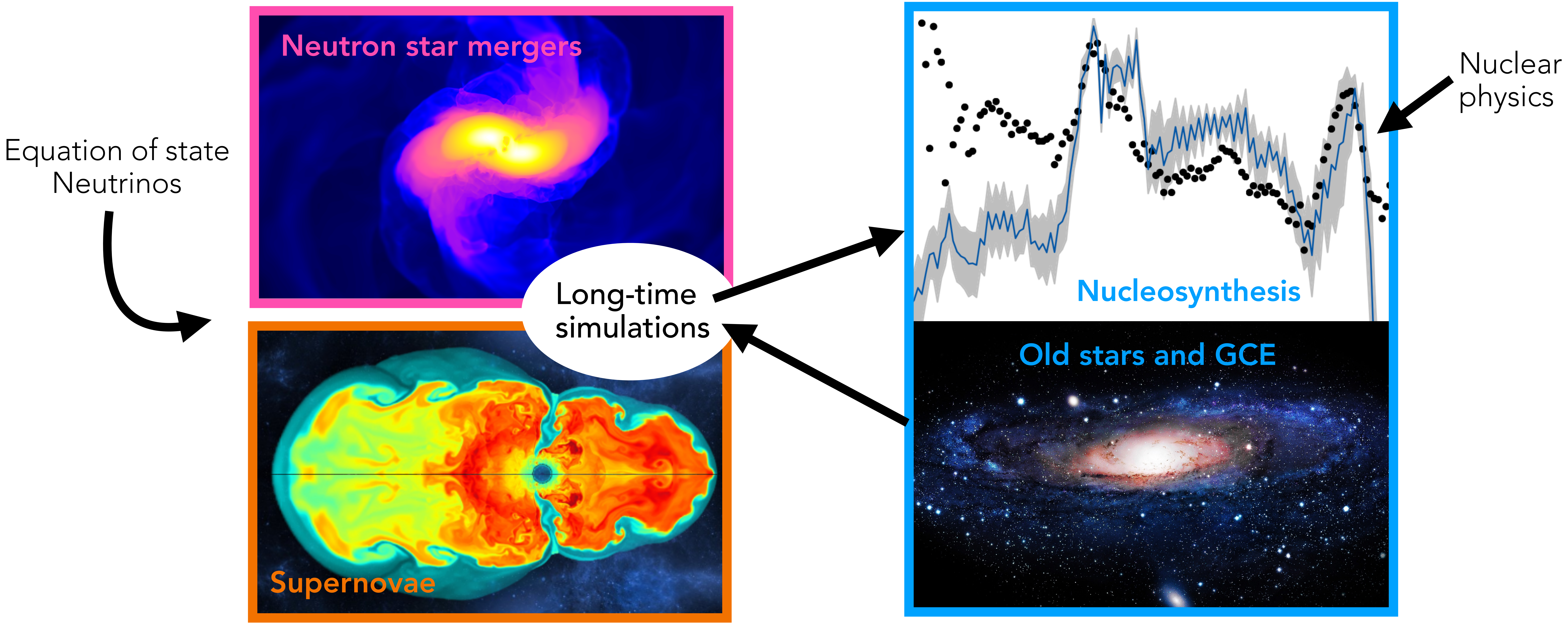


Côté et al. ApJ (2019)



Matteucci et al. MNRAS (2014), Côté et al. ApJ (2019), Molero et al. MNRAS (2021)

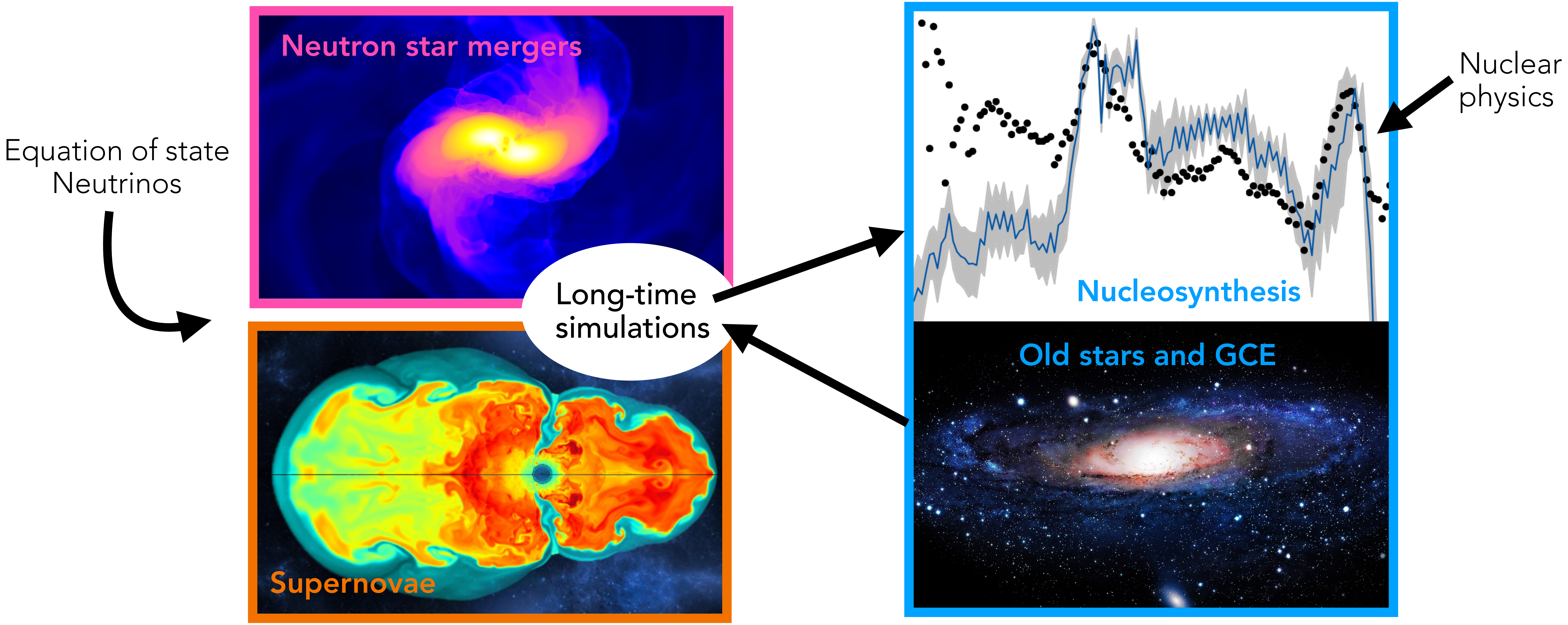
R-process: from simulations to observations



R-process: from simulations to observations

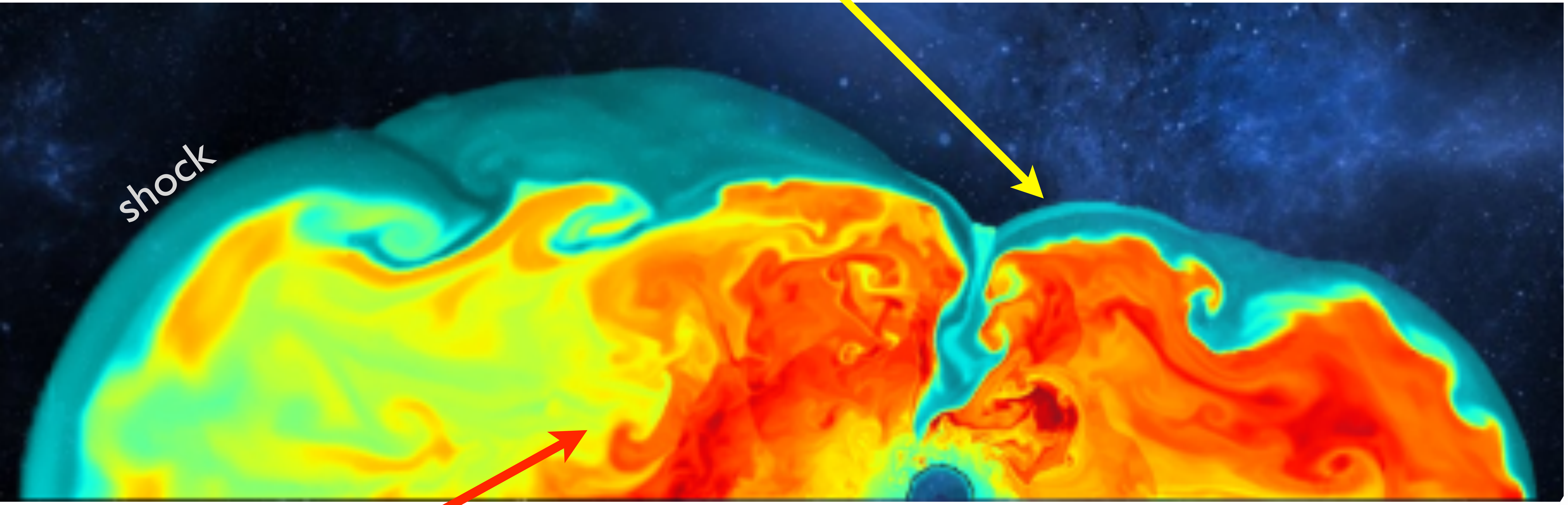
WinNet

https://github.com/nuc-astro Reichert et al. 2023



Supernova nucleosynthesis

Explosive nucleosynthesis: O, Mg, Si, S, Ca, Ti, Fe
shock wave heats falling matter

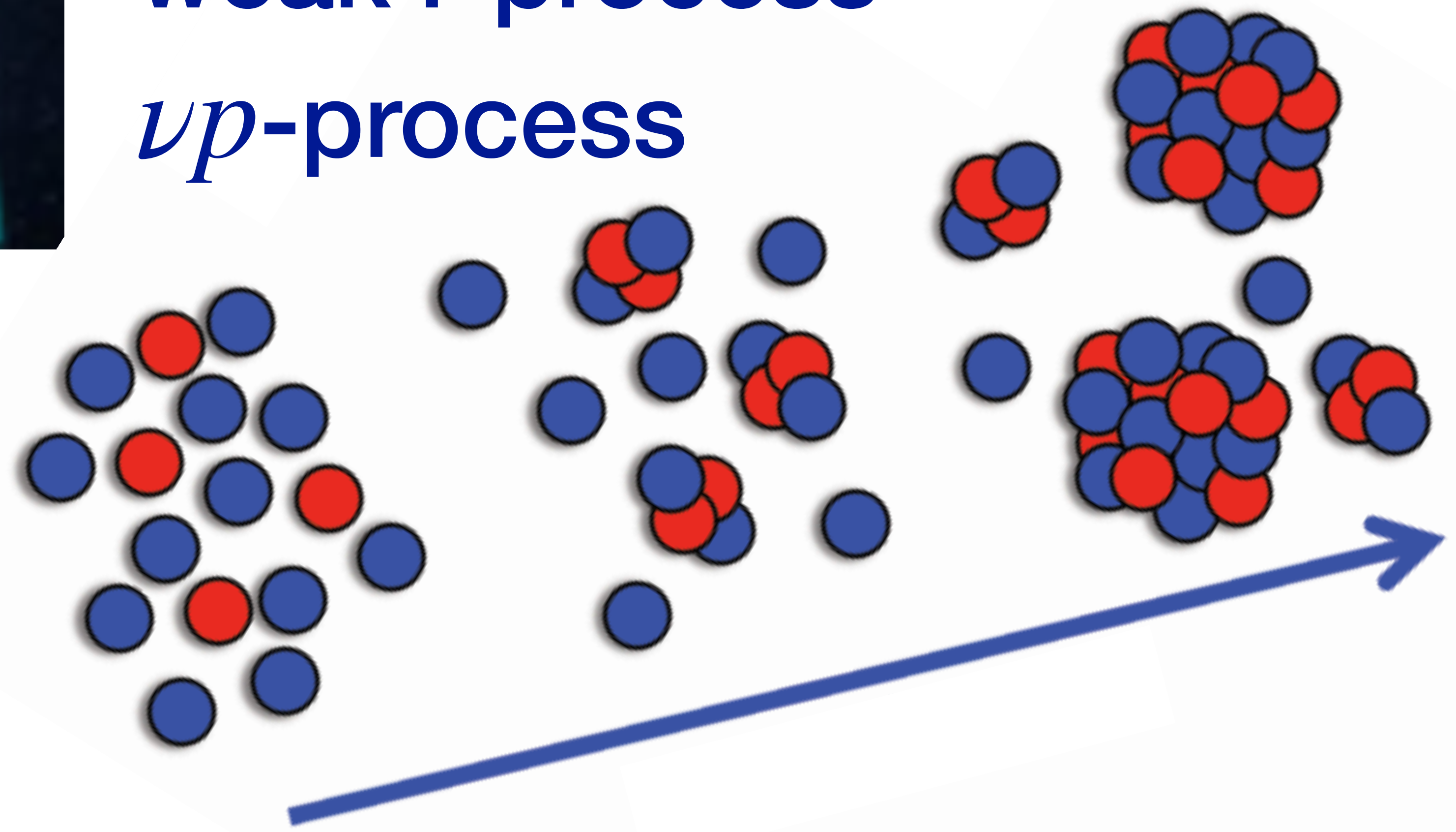


neutrino-driven ejecta

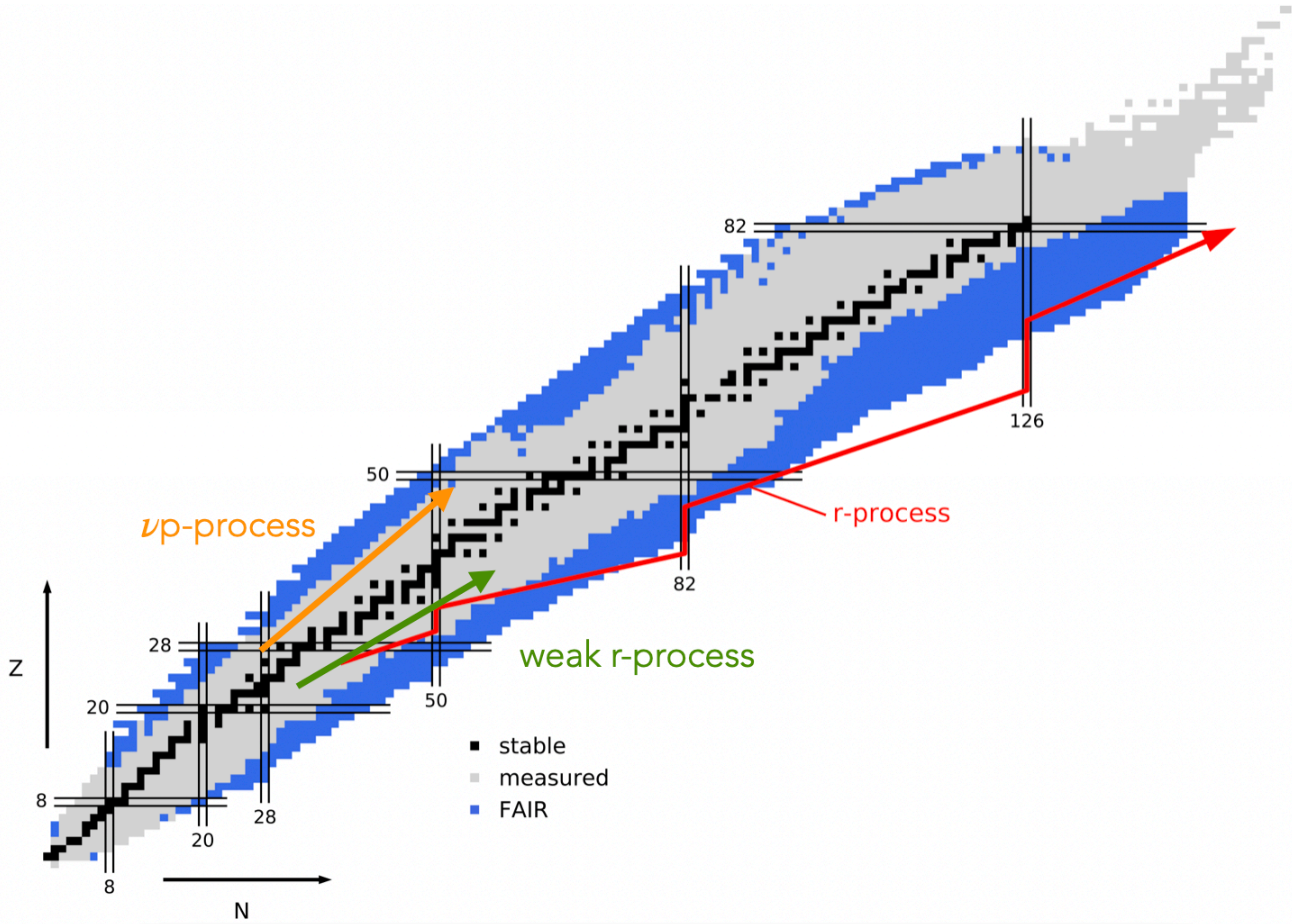
Nuclear statistical equilibrium (NSE)

charged particle reactions
α-process

r-process
weak r-process
νp-process



Supernova nucleosynthesis



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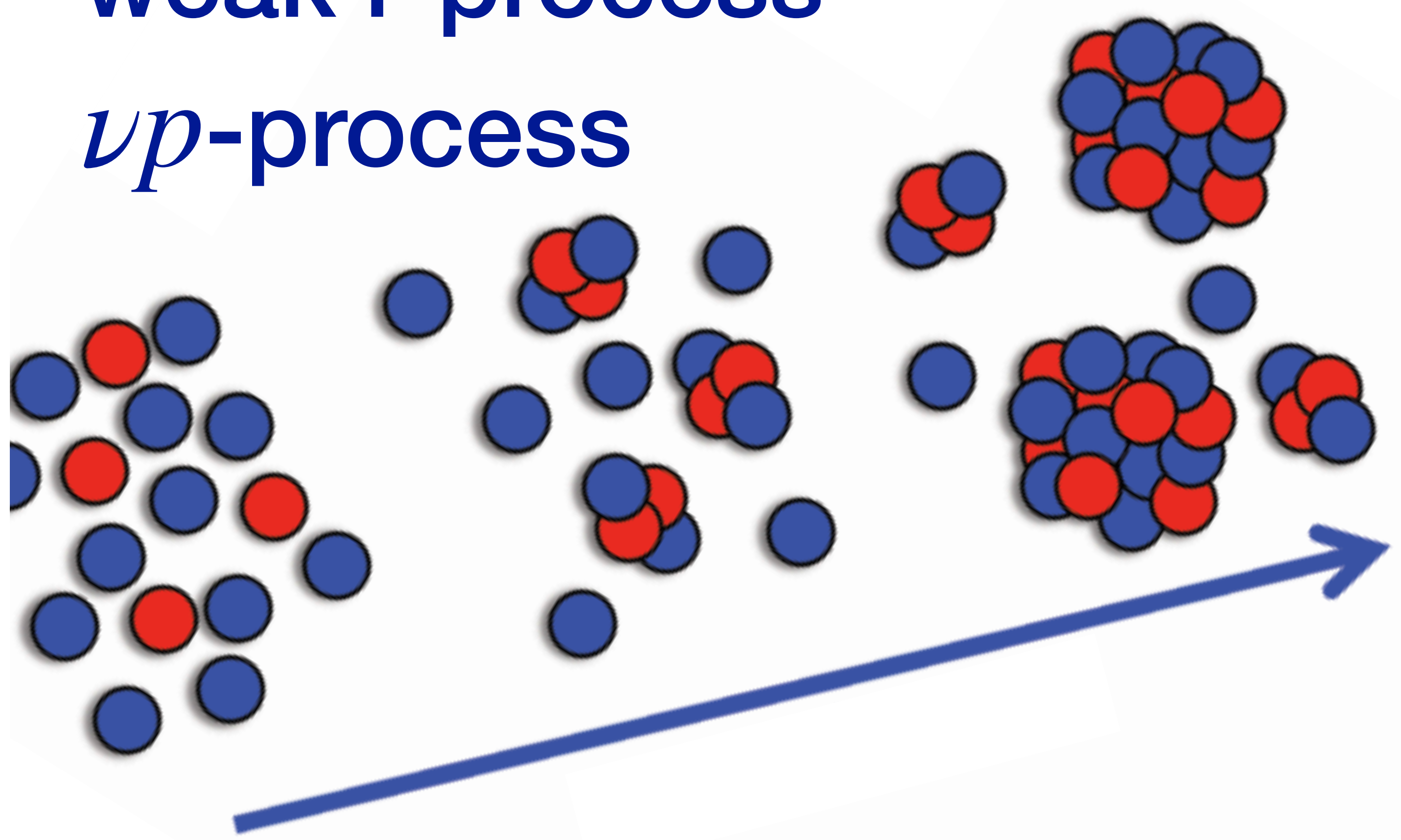
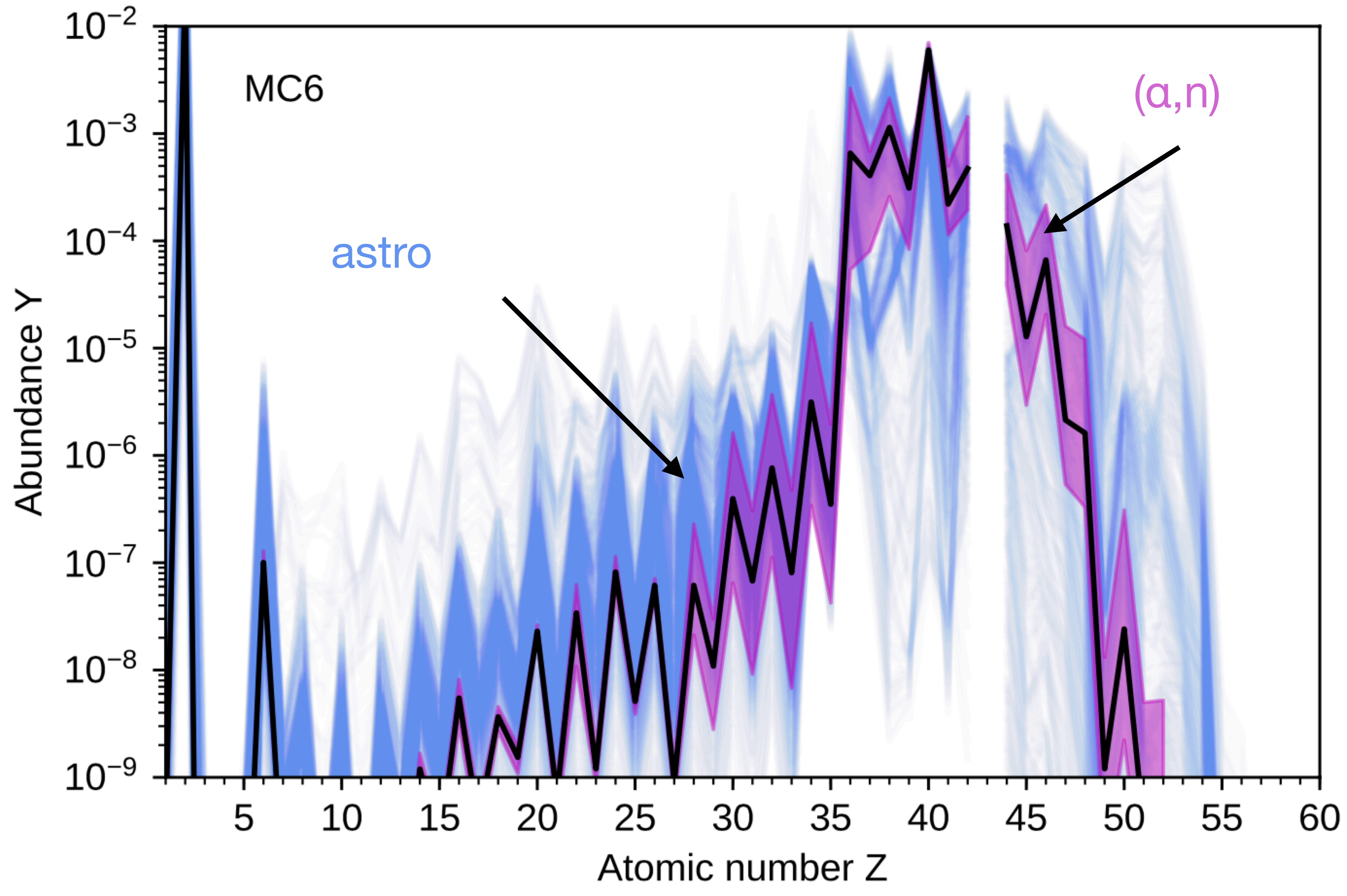


Core-collapse supernova: weak r-process

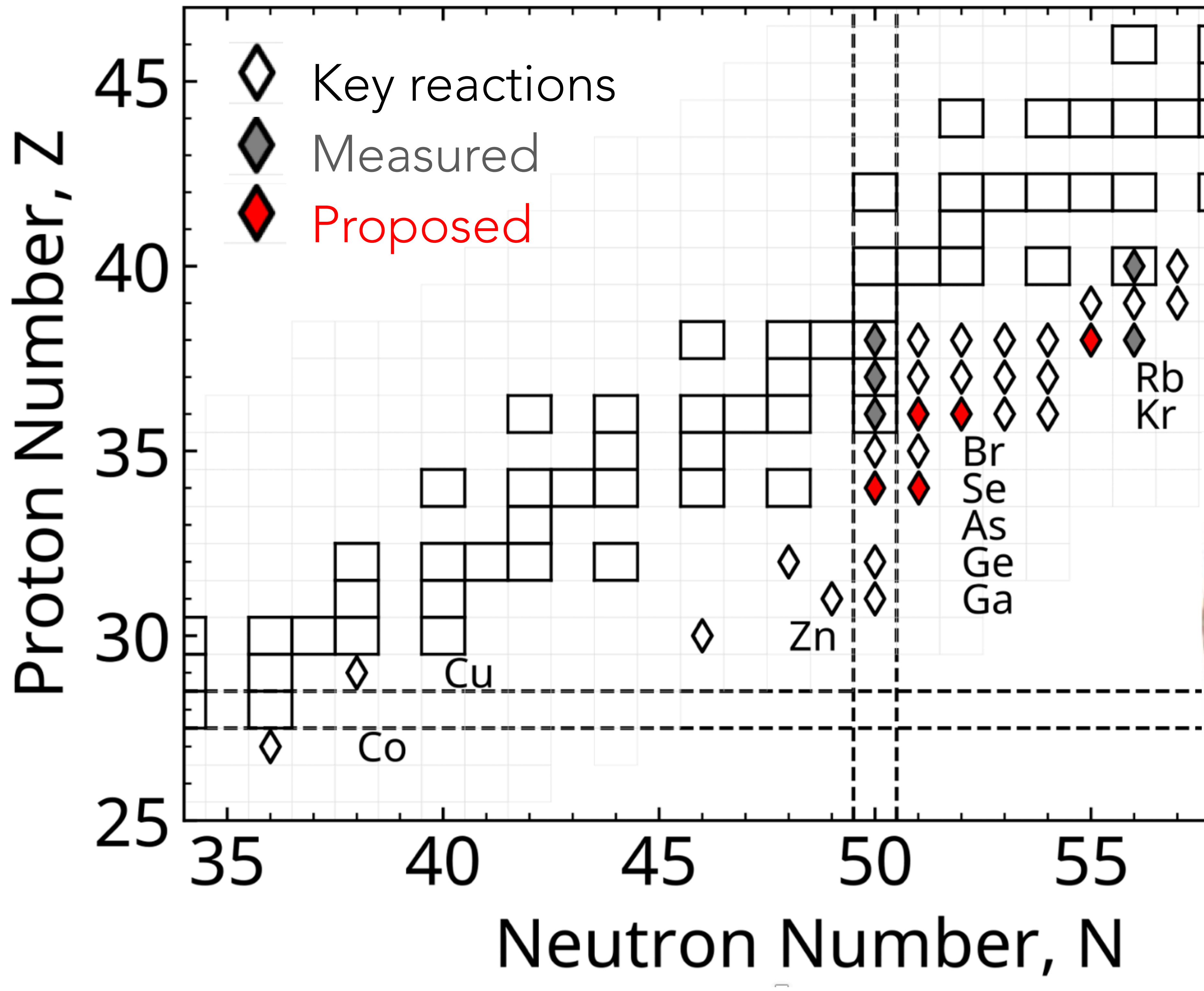
Neutrino-driven supernovae: elements up to Ag
Combine astrophysics and nuclear physics uncertainties
Motivation and support for experiments at NSCL, ANL, TRIUMF, ATOMKI

IReNA JINA-CEE NSF Physics Frontiers Center

Bliss et al. JPG (2017), Bliss et al. ApJ (2018), Bliss et al. PRC (2020)



Psaltis et al. ApJ (2022), Psaltis et al. ApJ (2024)

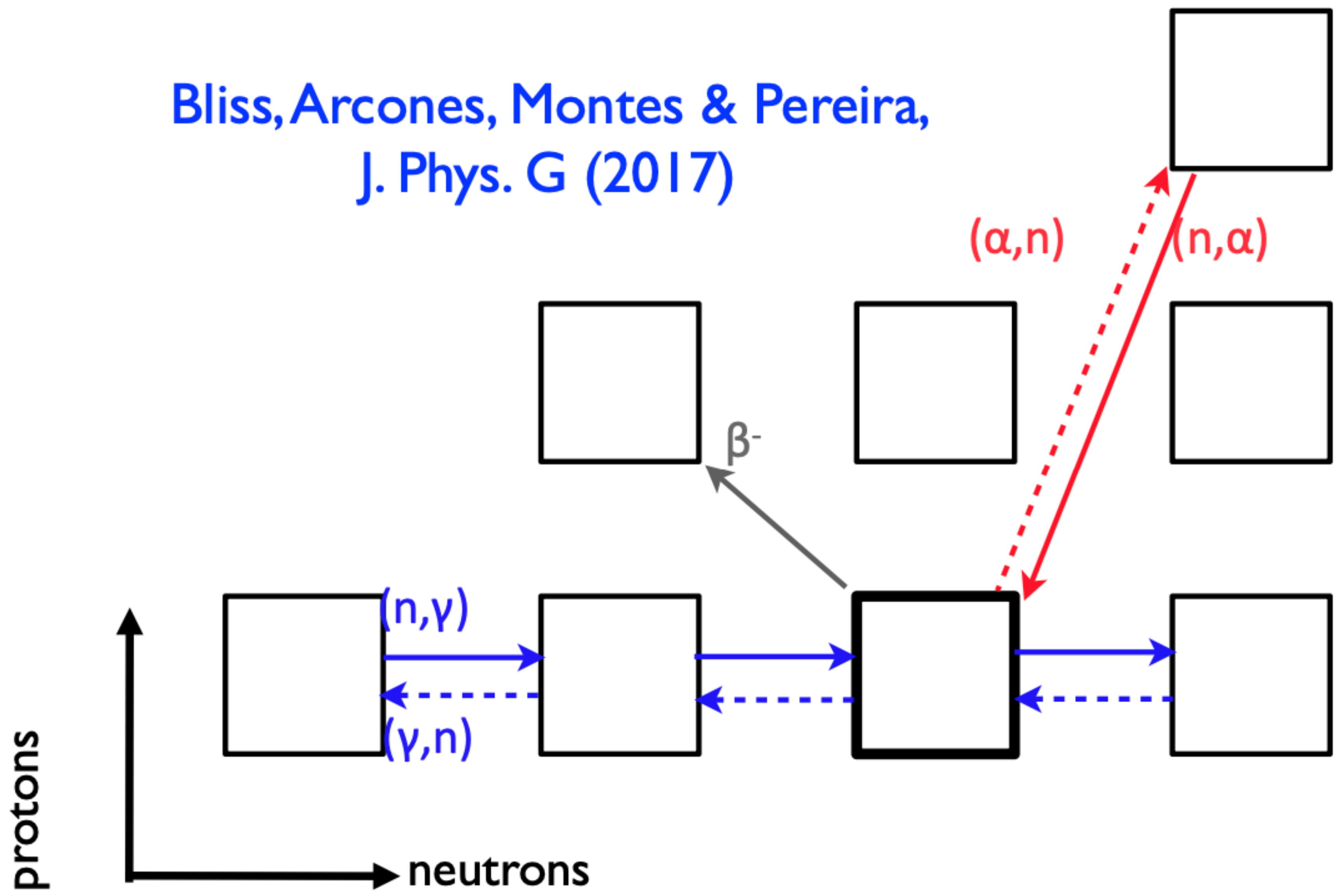


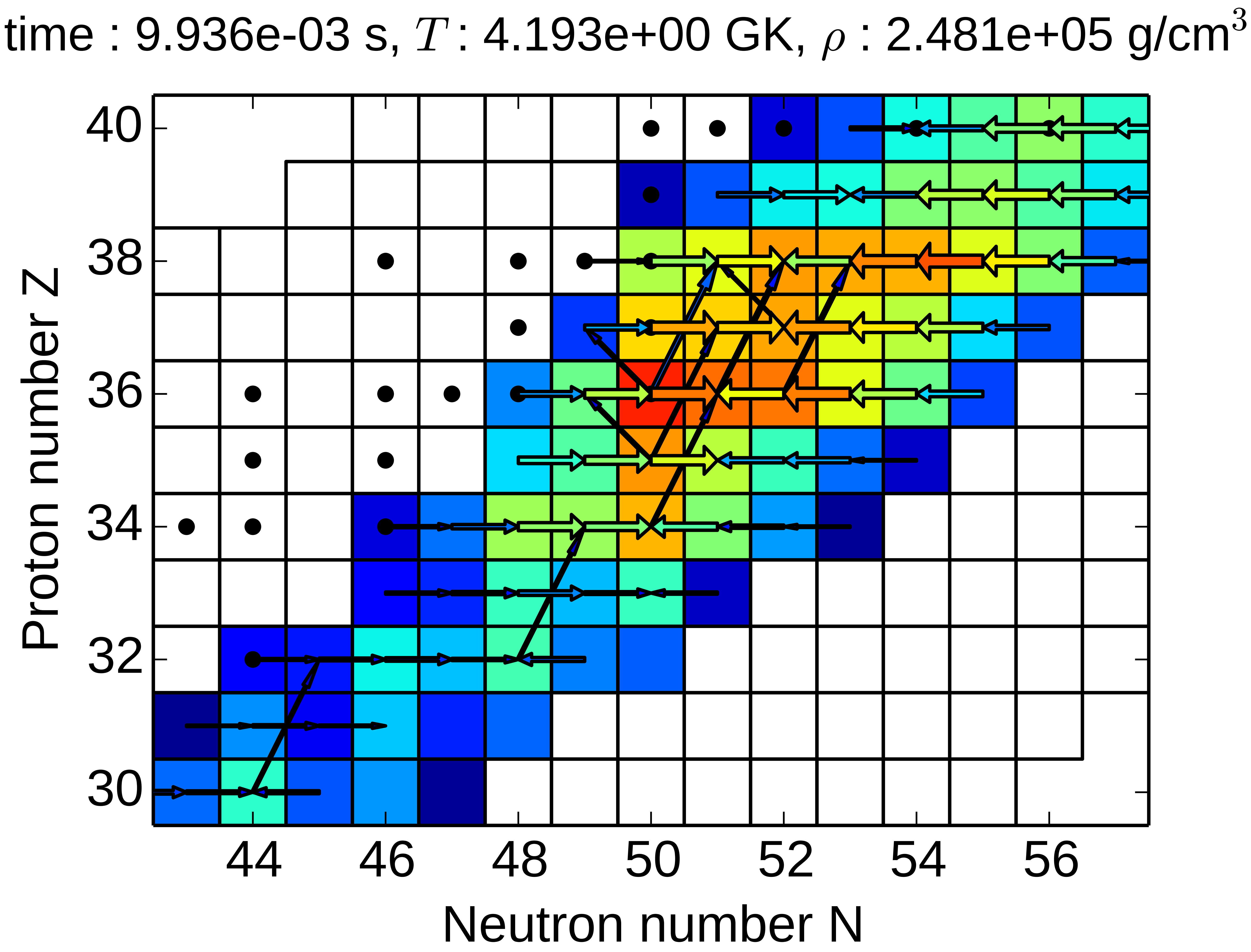
Nuclear physics uncertainty

Path close to stability:

- masses and beta decays known
- beta decays slow
- (α,n) reactions move matter to higher Z

Bliss, Arcones, Montes & Pereira, J. Phys. G (2017)

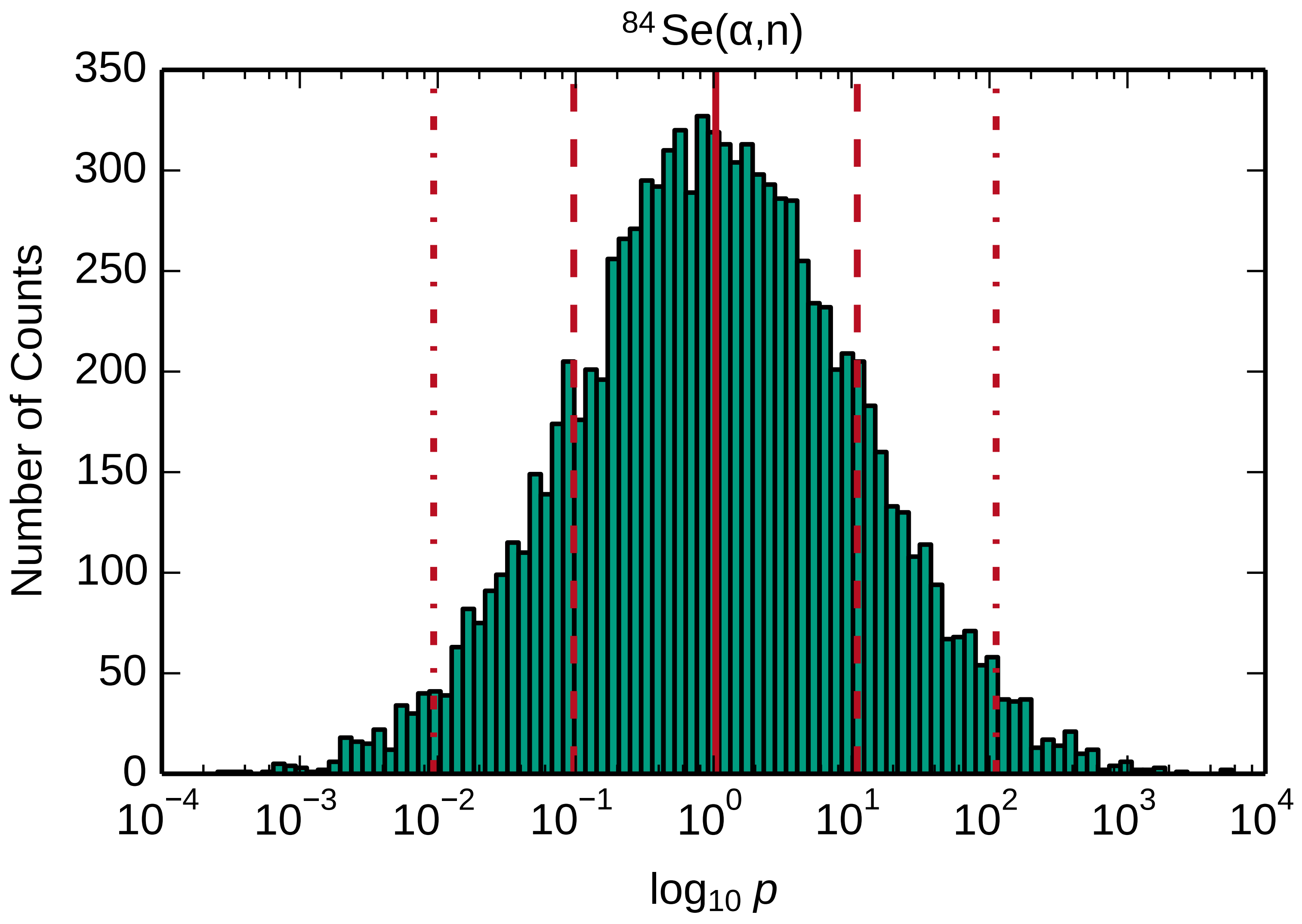


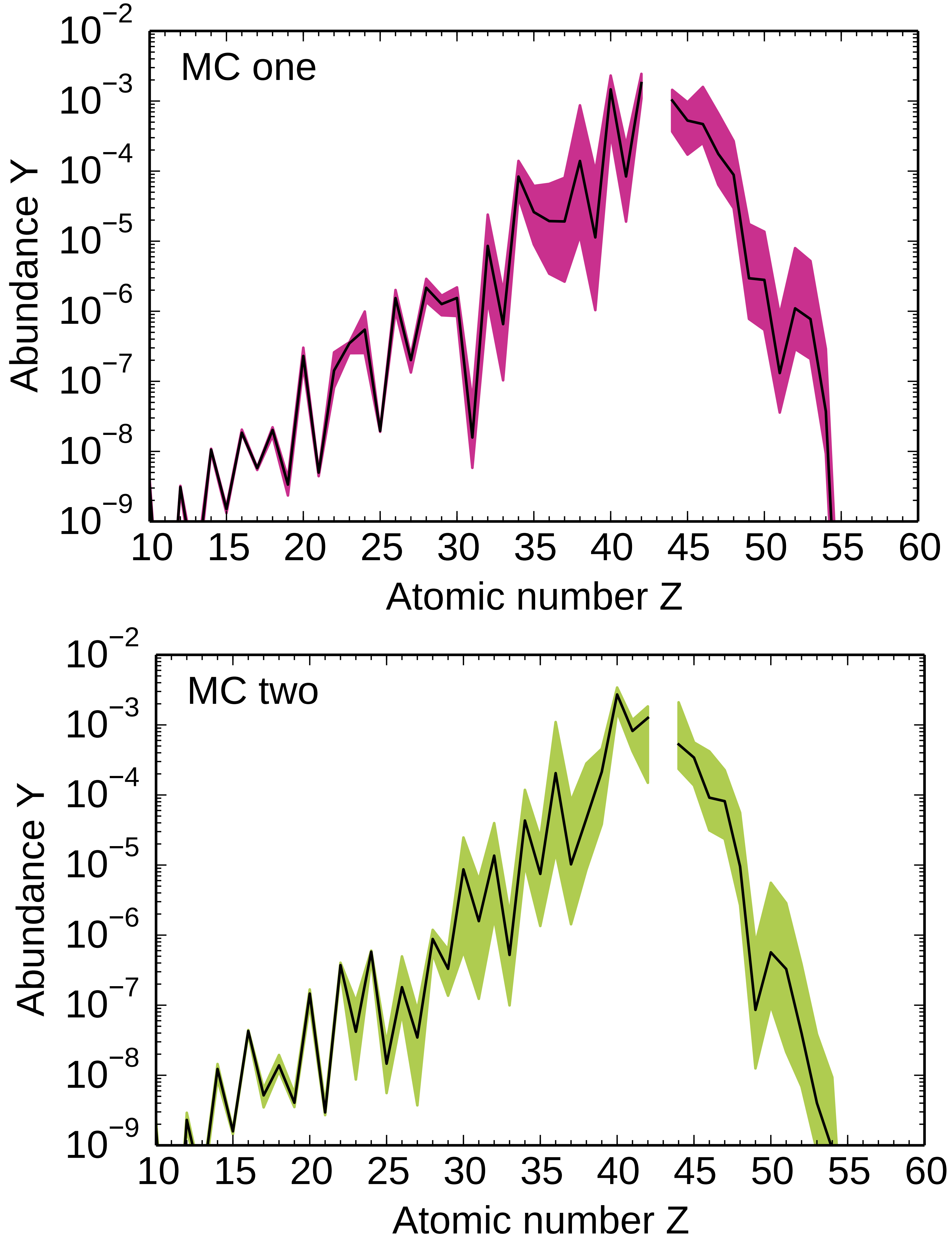


Sensitivity study

Independently vary each (α,n) reaction rate between Fe and Rh by a random factor

Include theoretical and experimental uncertainties
→ log-normal distributed rates (μ = 0, σ = 2.3)





36 representative trajectories
10 000 Monte Carlo runs

Bliss et al., PRC (2020)

Sensitivity study: key reactions

Bliss et al., PRC (2020), Psaltis et al., ApJ (2022)

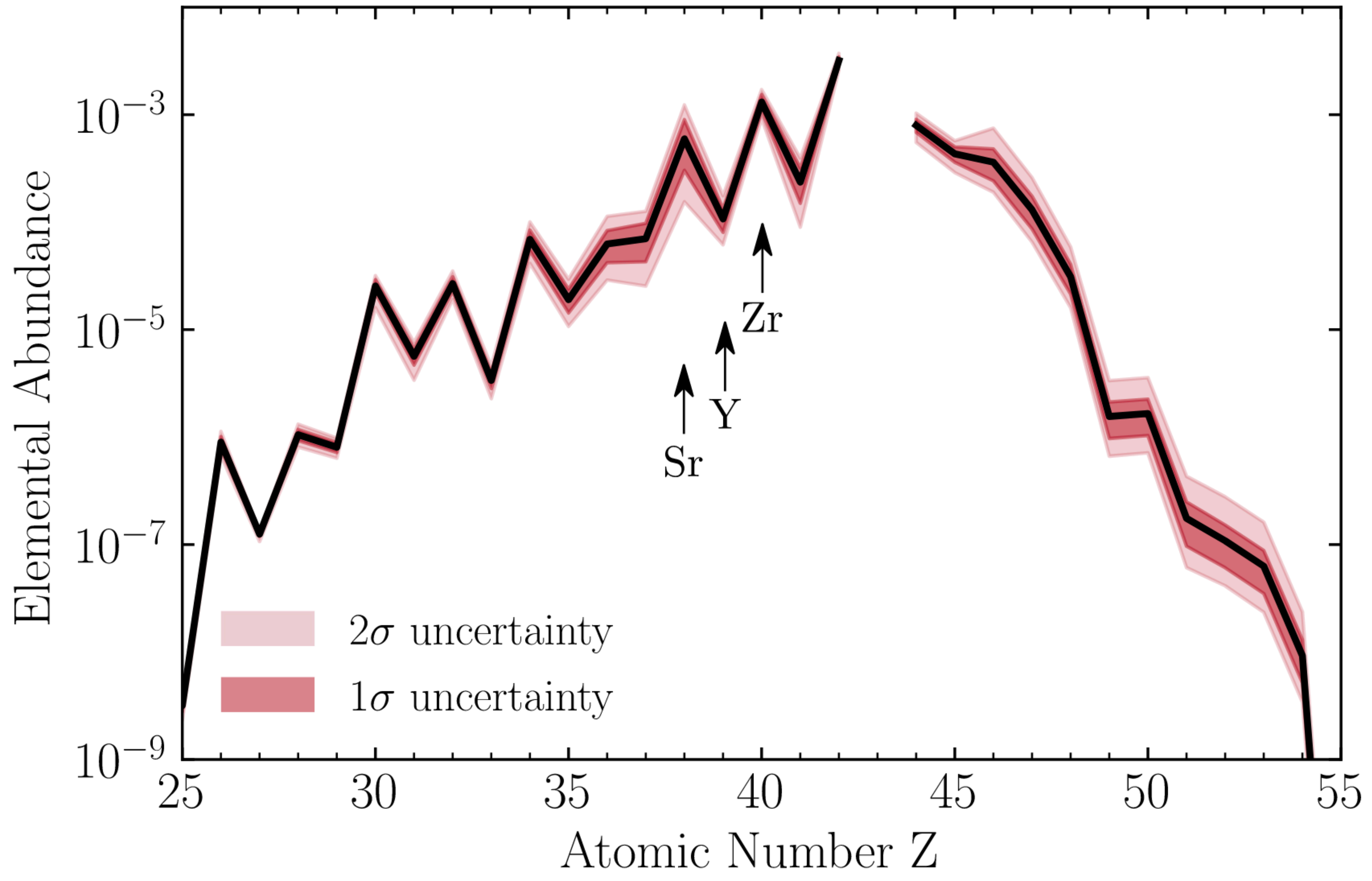
Key reactions ⇒ large correlation + significant impact on abundance for several astro conditions

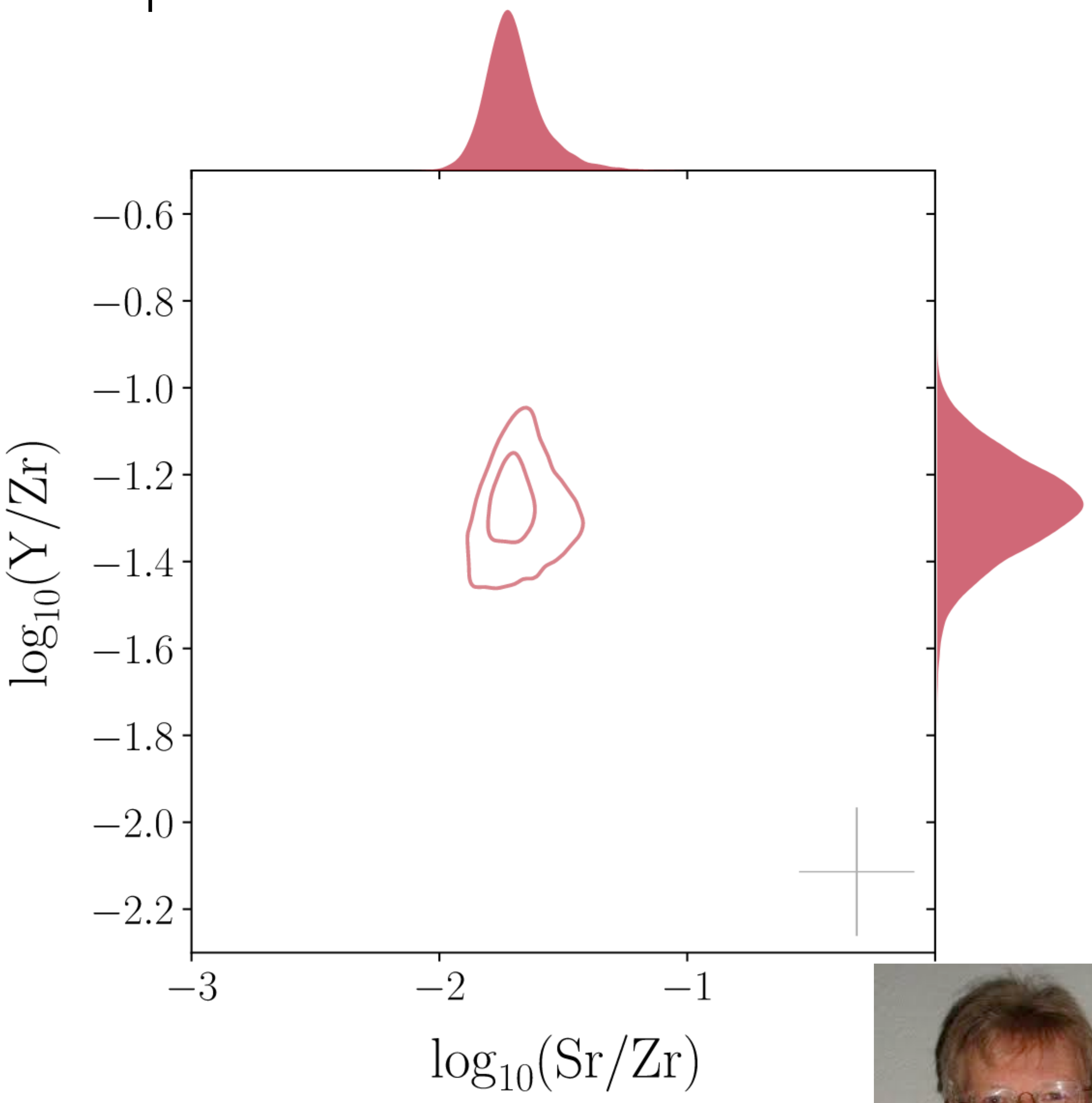
Reaction	Z	MC tracers
^{59}Fe$(\alpha, n)^{62}$Ni	39 – 42, 45	34, 36
^{68}Fe$(\alpha, n)^{71}$Ni	36, 37	3
^{63}Co$(\alpha, n)^{66}$Cu	39–42, 45	20, 34, 36
^{71}Co$(\alpha, n)^{74}$Cu	36, 37	3
^{74}Ni$(\alpha, n)^{77}$Zn	36–42	2, 3, 17, 18, 32
^{76}Ni$(\alpha, n)^{79}$Zn	36–42	2, 3, 18, 32
^{67}Cu$(\alpha, n)^{70}$Ga	47	35
^{77}Cu$(\alpha, n)^{80}$Ga	37	3
^{72}Zn$(\alpha, n)^{75}$Ge	39–42	36
^{76}Zn$(\alpha, n)^{79}$Ge	36, 37–42	2, 3, 17, 18, 32
^{78}Zn$(\alpha, n)^{81}$Ge	36, 37–42	2, 3, 17, 18, 32
^{79}Zn$(\alpha, n)^{82}$Ge	36, 37–42	2, 3, 18, 32
^{80}Zn$(\alpha, n)^{83}$Ge	36, 37, 39–42	2, 3, 18, 32
^{81}Ga$(\alpha, n)^{84}$As	36, 38, 39, 41	17, 32
^{78}Ge$(\alpha, n)^{81}$Se	39–42	36
^{80}Ge$(\alpha, n)^{83}$Se	36–39, 42	28, 33, 36
^{82}Ge$(\alpha, n)^{85}$Se	36–39, 41	11, 17, 19, 27, 28, 33
^{83}As$(\alpha, n)^{86}$Br	36, 37, 41	11, 26, 27, 28, 33
^{84}Se$(\alpha, n)^{87}$Kr	36–42, 44, 45	2, 6, 7, 8, 9, 10, 11, 18, 19, 20, 22, 23, 24, 26, 27, 28, 29, 30, 31, 33, 34, 36
^{85}Se$(\alpha, n)^{88}$Kr	36–42, 44, 45	2, 6, 7, 8, 9, 10, 11, 18, 19, 22, 23, 24, 26, 27, 28, 29, 30, 31
^{85}Br$(\alpha, n)^{88}$Rb	37–39	6, 7, 8, 9, 10, 22, 23, 24, 26, 28, 29, 30, 31
^{87}Br$(\alpha, n)^{90}$Rb	37, 39	6, 9, 10, 29, 31
^{88}Br$(\alpha, n)^{91}$Rb	39	26
^{86}Kr$(\alpha, n)^{89}$Sr	38–42, 44, 45, 47	4, 5, 7, 8, 13, 14, 15, 16, 20, 24, 25, 33, 34, 35

Comparison to observations

Psaltis et al., ApJ (2022)

Abundance with uncertainties for several astro conditions $\longrightarrow$ compare abundance ratios



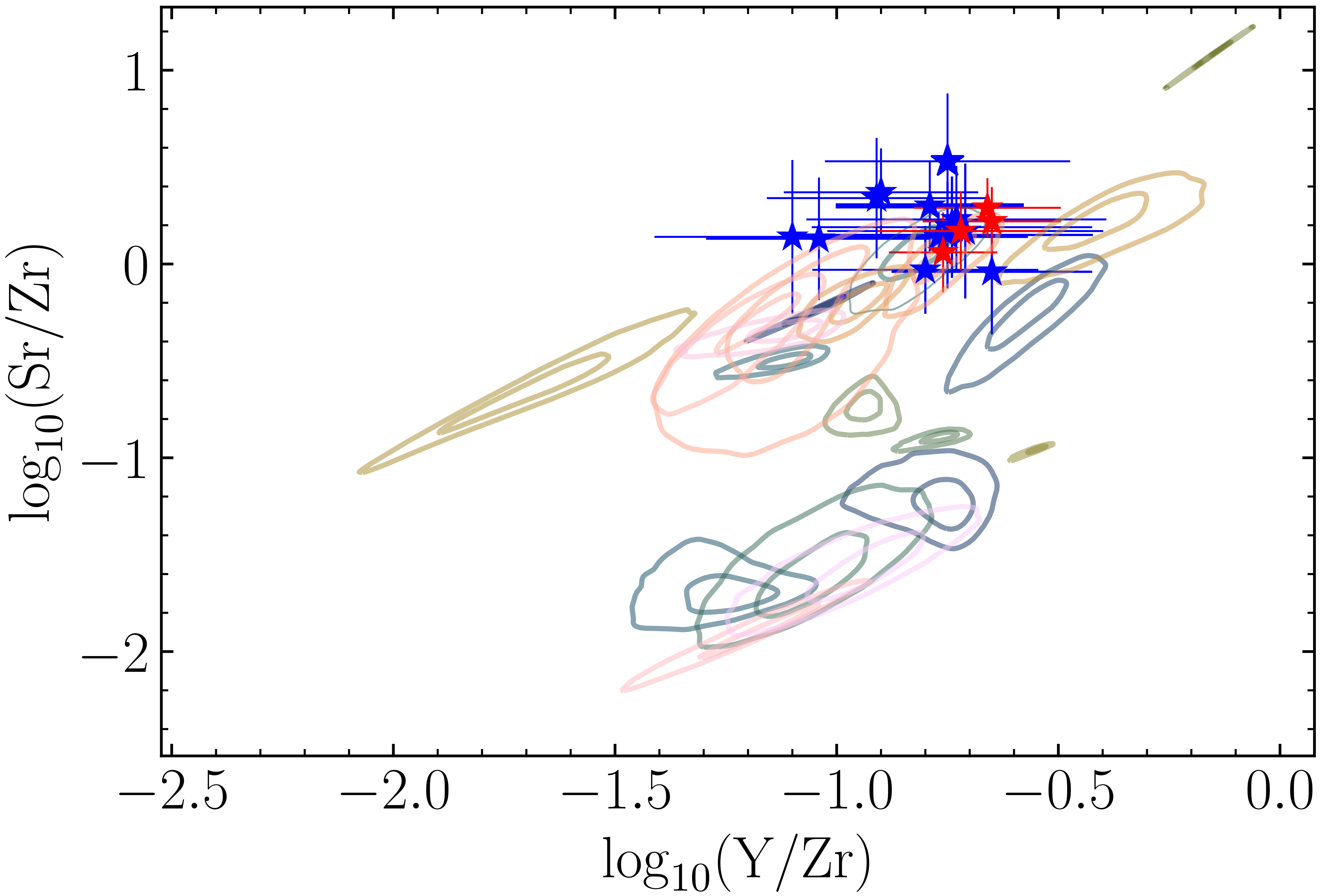


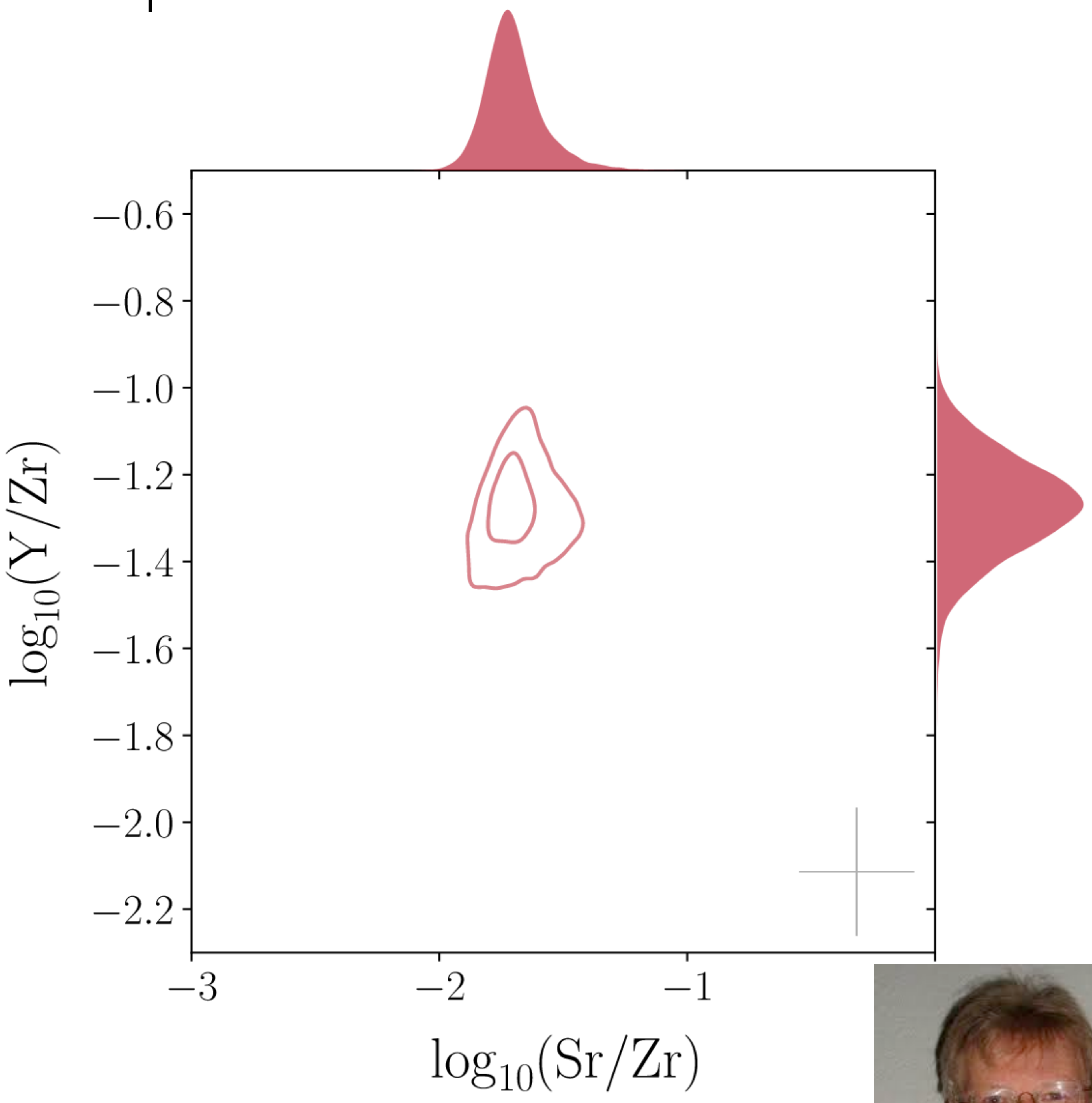
Based on optical potentials from Mohr et al., ADNDT (2021)

Comparison to observations

Psaltis et al., ApJ (2022)

Abundance with uncertainties for several astro conditions ⟶ compare abundance ratios





Based on optical potentials from Mohr et al., ADNDT (2021)

What has been measured so far?

- ^{86}Kr(α, n), ^{96}Zr(α, n) and ^{100}Mo(α, n) at ATOMKI

  G.G. Kiss et al., Astrophys. J **908**, 202 (2021) • T.N. Szegedi et al., Phys. Rev. C **104**, 035804 (2021)

György Gyürky
Poster: Sándor Kovács

- ^{75}Ga(α, n), ^{85,86}Kr(α, xn), ^{85}Br(α, xn) at NSCL/FRIB (HabaNERO/SECAR)

  F. Montes, J. Pereira et al.

- ^{86}Kr(α, xn), ^{87}Rb(α, xn), ^{88}Sr(α, xn), ^{100}Mo(α, xn) at Argonne (MUSIC)

  M. L. Avila, C. Fougères et al.

  W. J. Ong et al., Phys. Rev. C **105**, 055803 (2022)

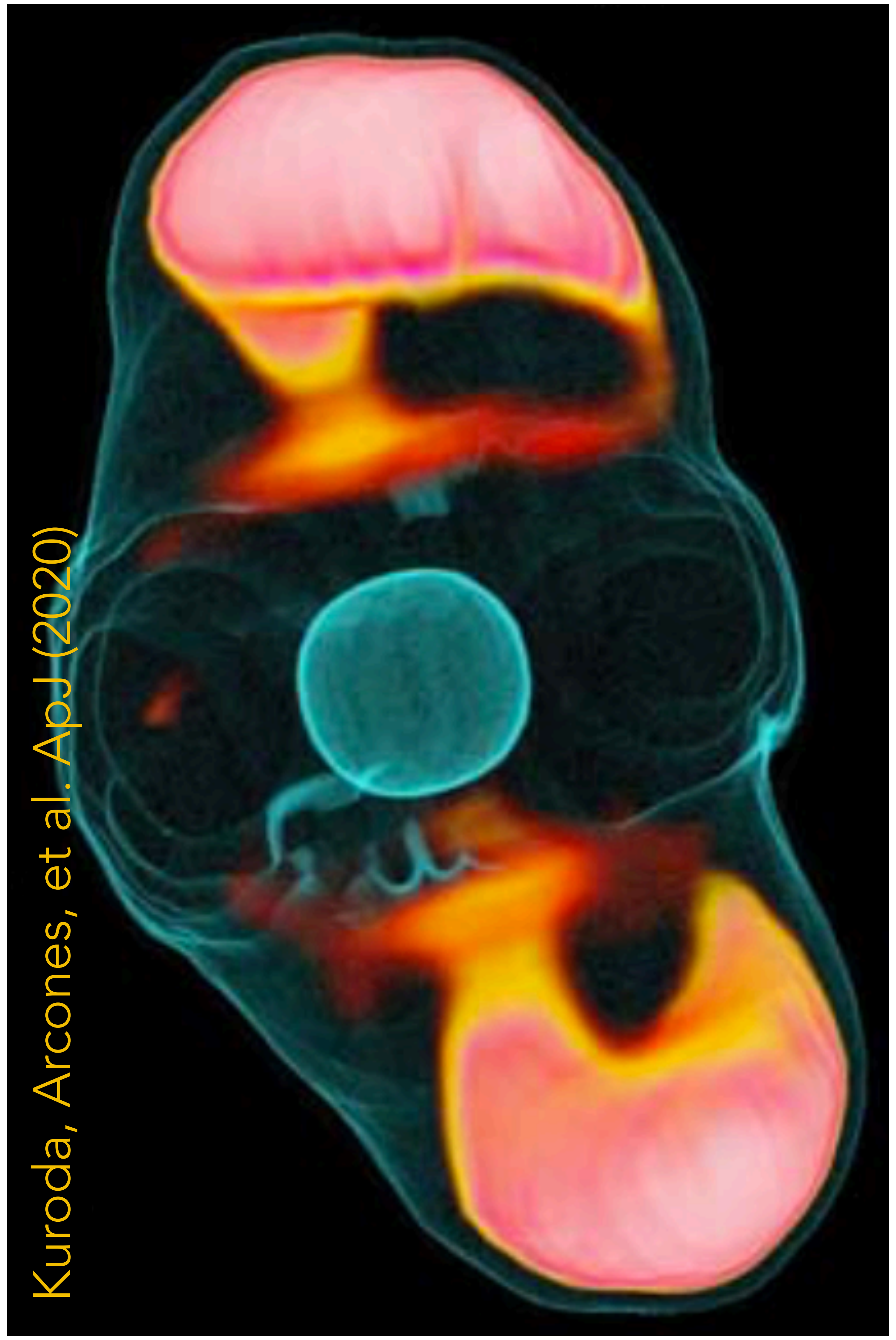
- ^{86}Kr(α, n) and ^{94}Sr(α, n) at TRIUMF (EMMA)

  C. Aa. Diget, A. M. Laird, M. Williams et al.

  C. Angus *et al.*, EPJ Web of Conferences, NPA-X (2023)

r-process in supernovae?

- Neutrino-driven supernovae: elements up to Ag
- Magneto-rotational supernovae: elements up to U and Th?



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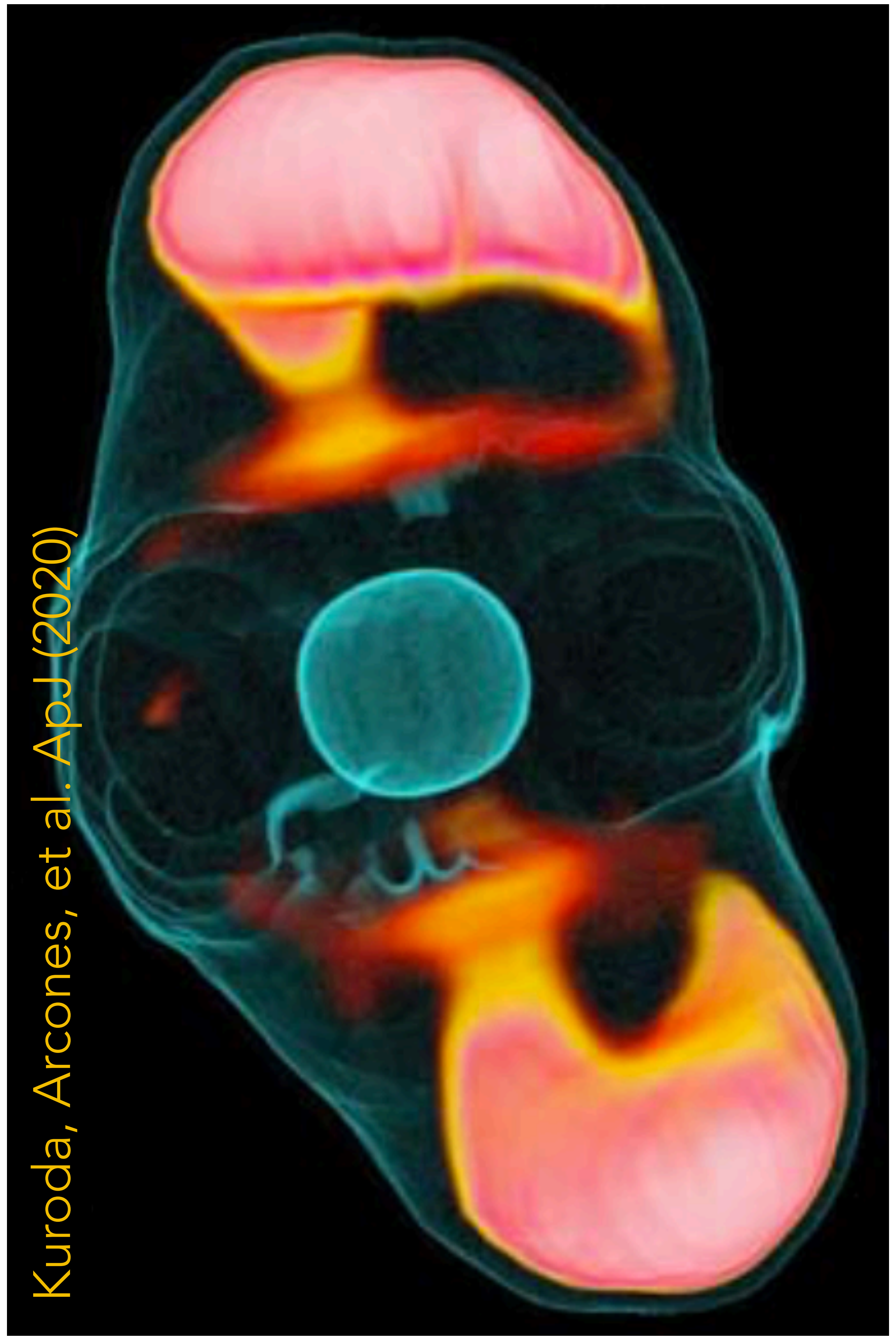

Neutron-rich matter ejected by magnetic field (Cameron 2003, Nishimura et al. 2006)

2D and 3D + parametric neutrino treatment

Winteler et al. 2012, Nishimura et al. 2015, 2017, Mösta et al. 2018

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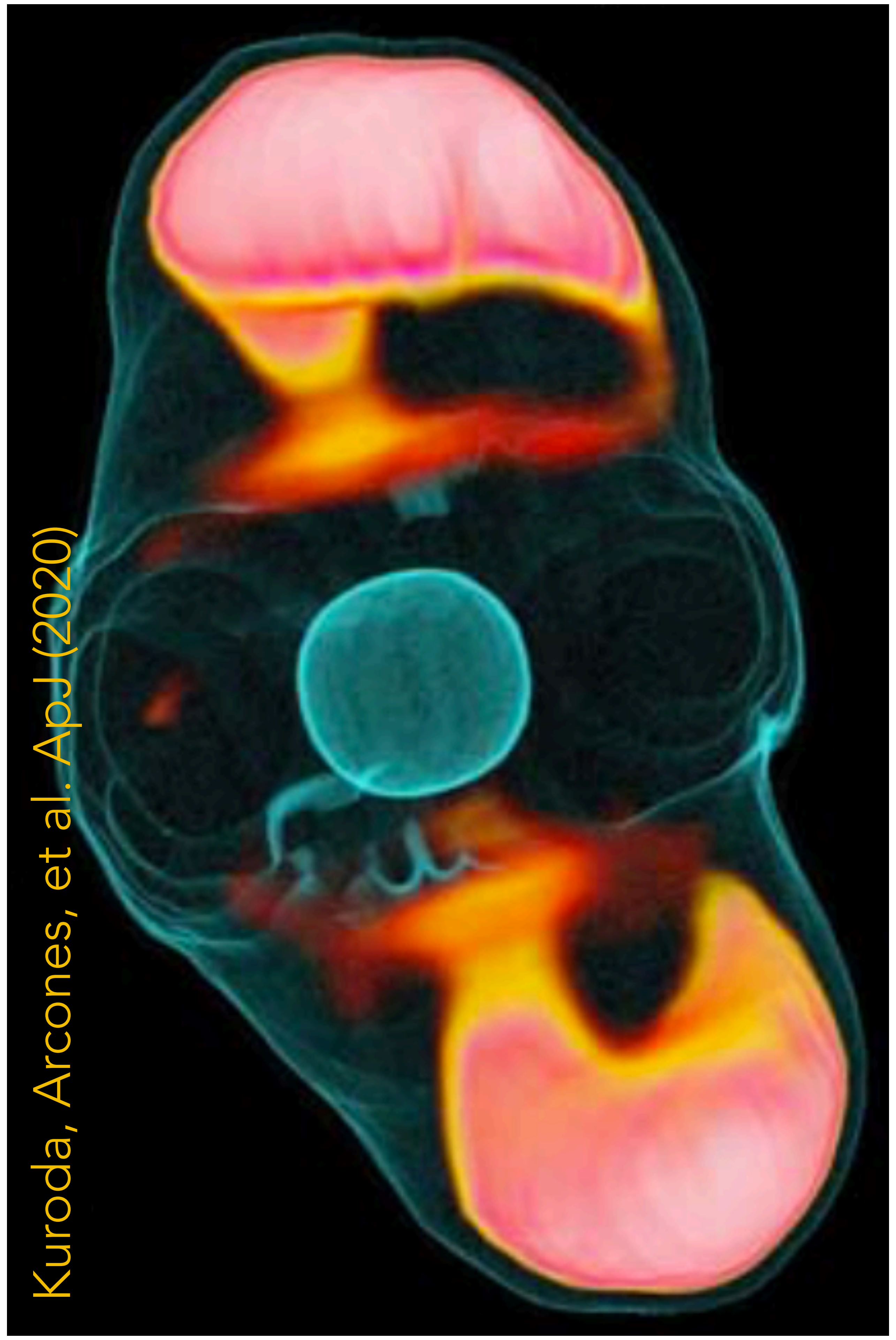

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First simulations of explosions with magnetic fields and detailed neutrino transport (Obergaulinger & Aloy 2017), and their nucleosynthesis (Reichert et al. ApJ 2021, Reichert et al. MNRAS 2023)

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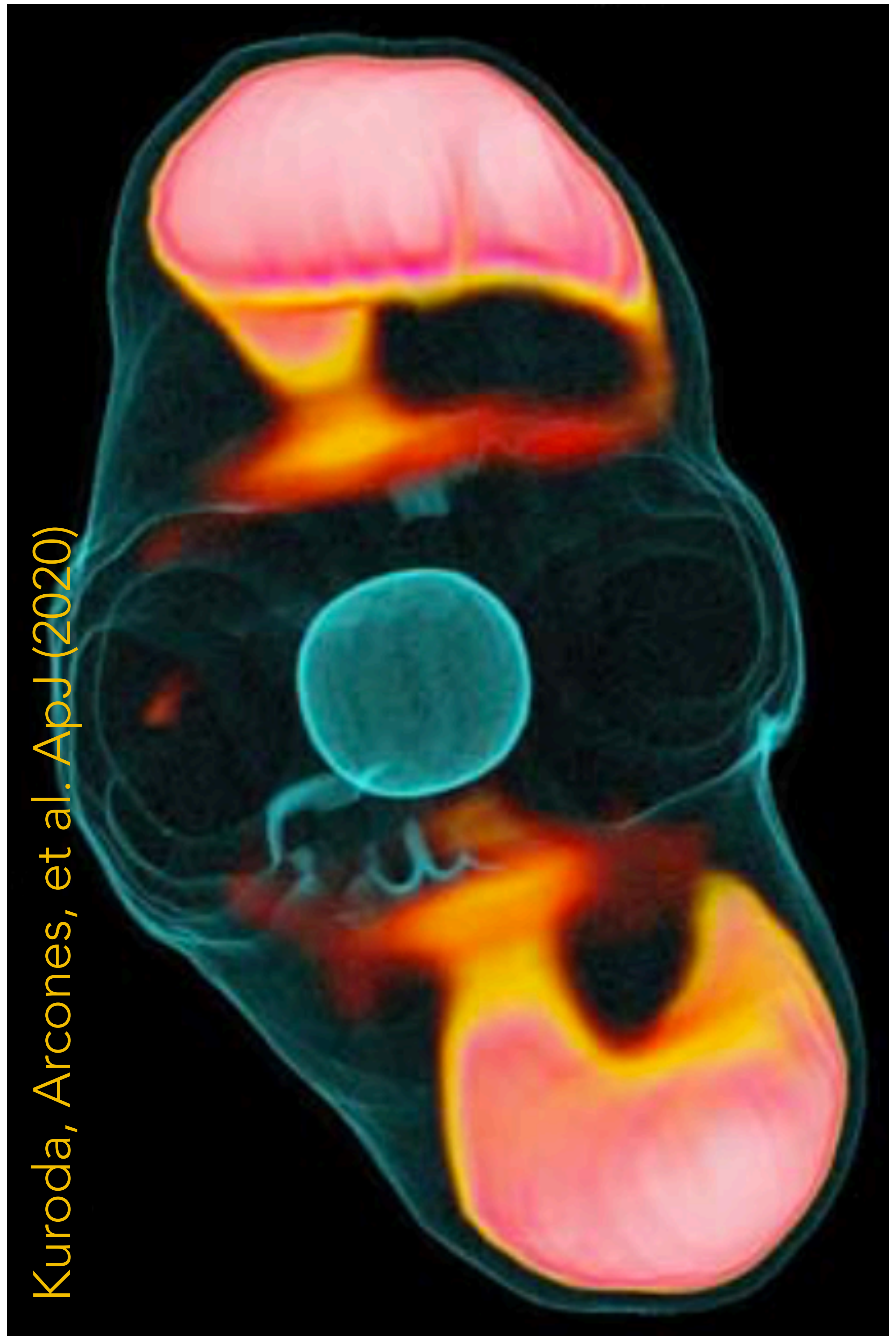

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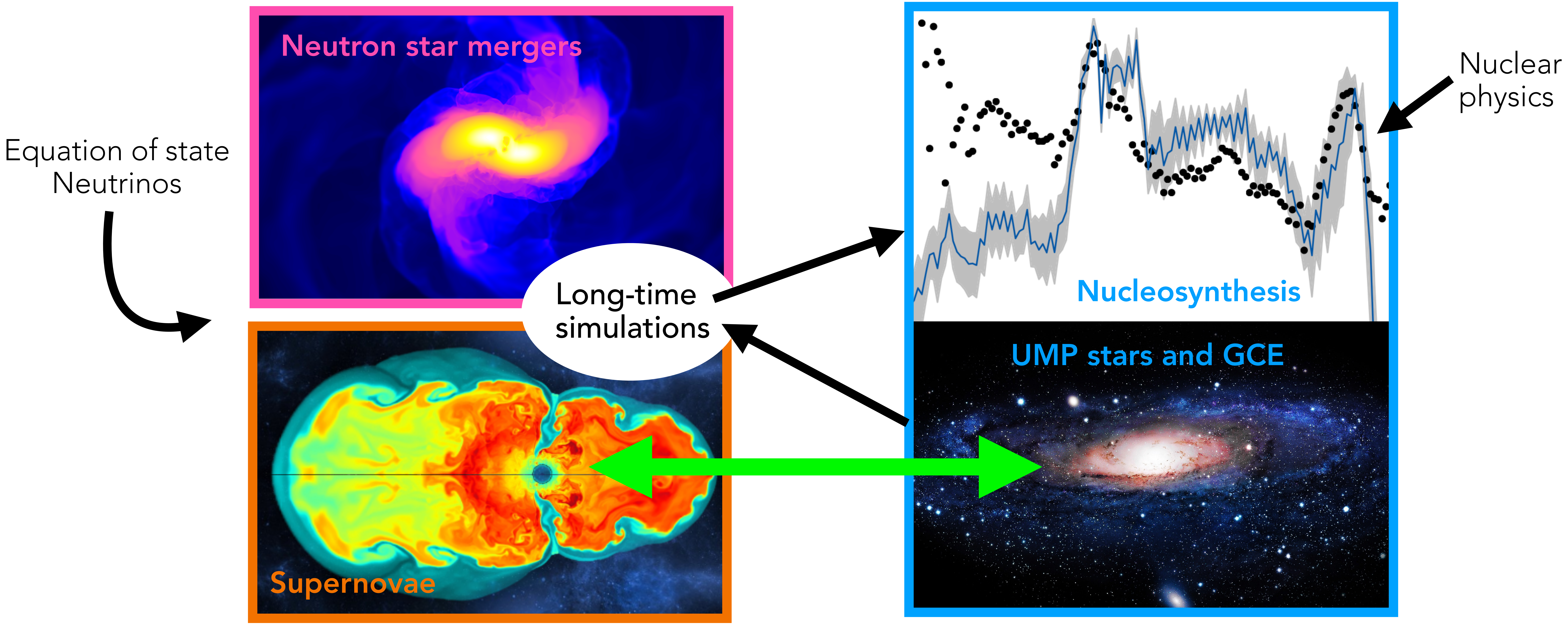
Open questions

- Long-time evolution:
  Magnetar (neutron star) vs. Collapsar (black hole): **r-process possible?**
- Impact of magnetic field strength and morphology on nucleosynthesis

Reichert et al. MNRAS (2024)

R-process: from simulations to observations

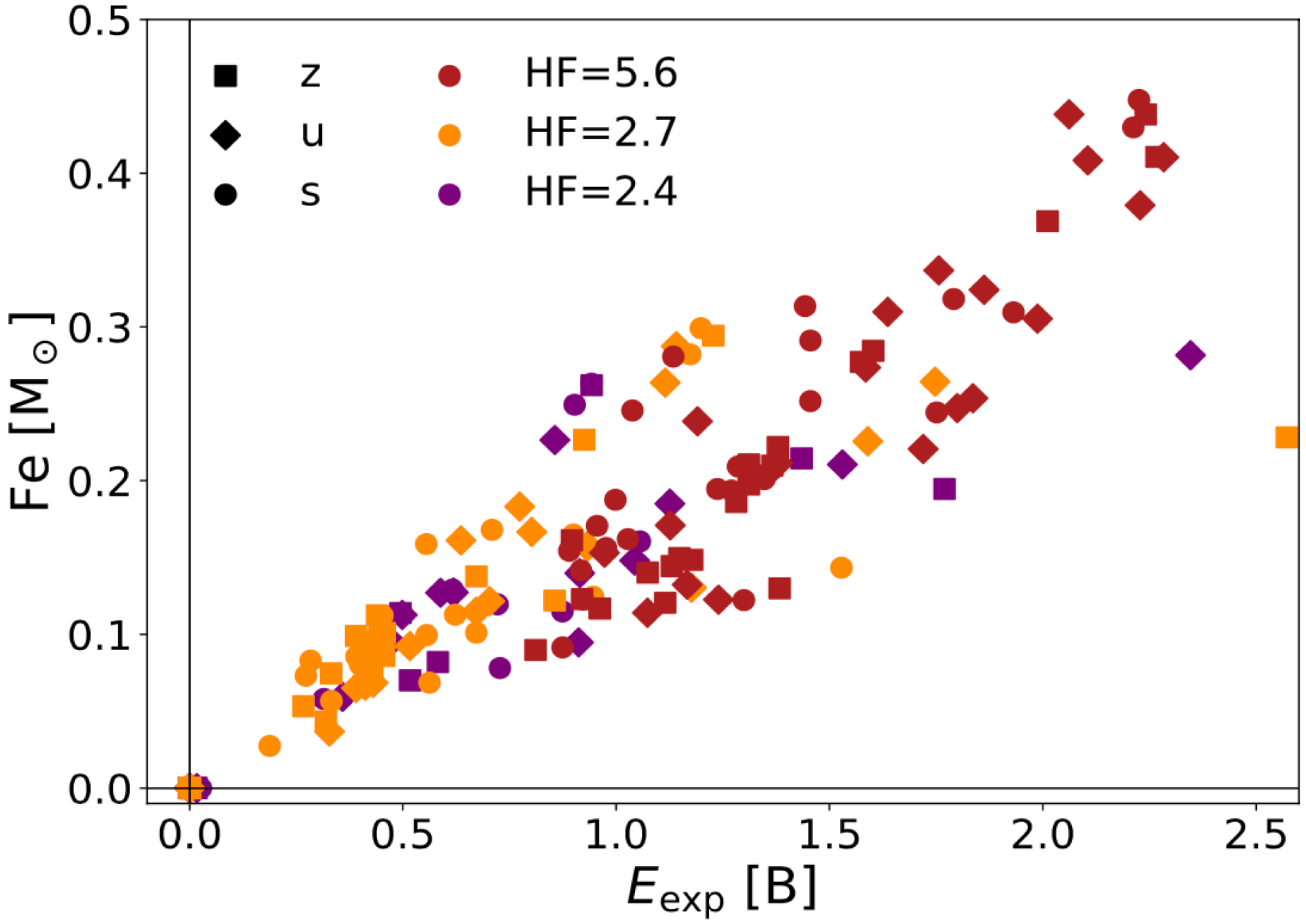


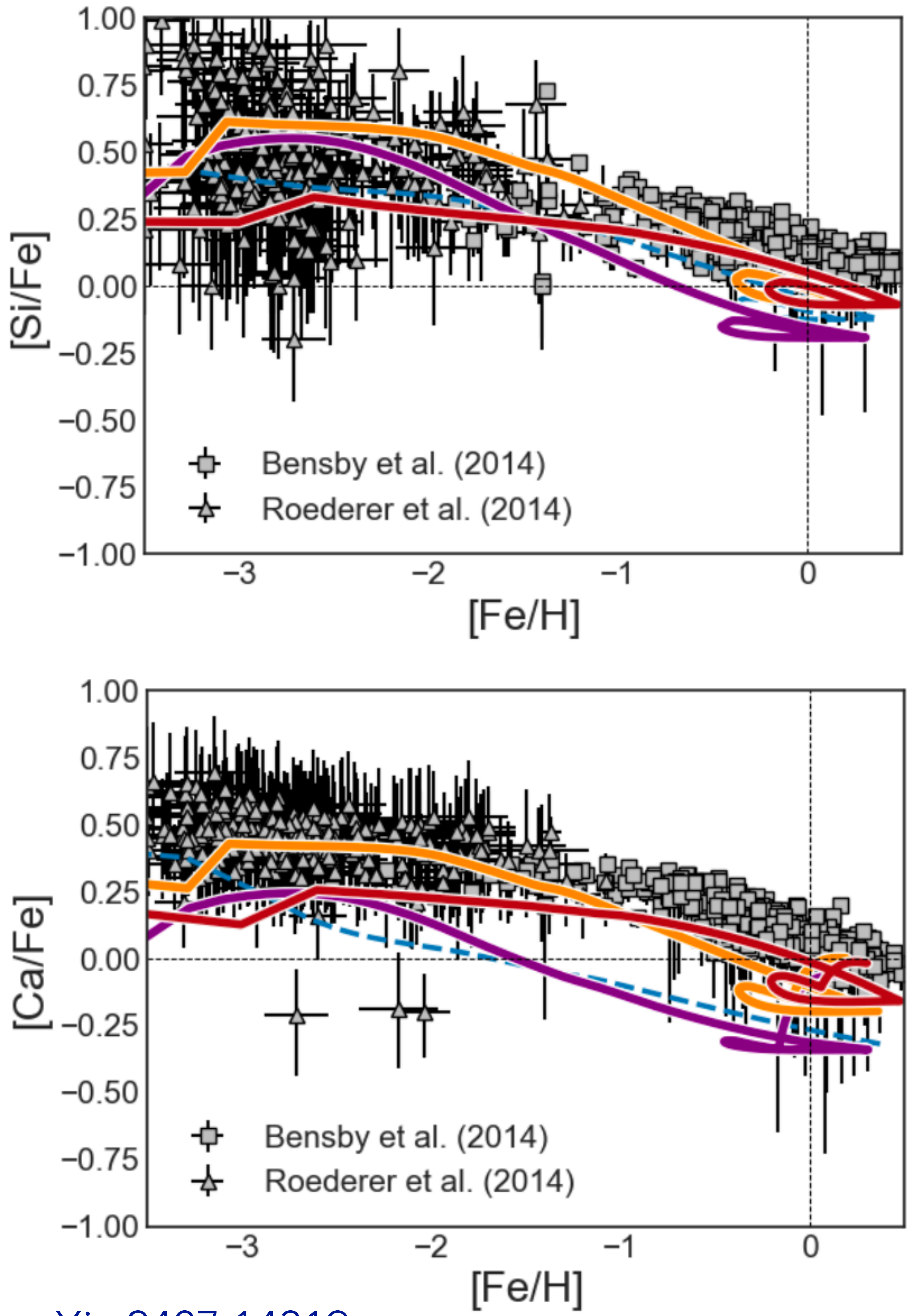
Core-collapse supernova yields for galactic chemical evolution (GCE)

Reduced alpha-network within simulations (Navó et al. 2023)

189 simulations,1D + accurate neutrino transport + neutrino heating

⟶propagate uncertainties from supernova models to GCE





Jost, Molero et al. submitted, arXiv:2407:14319

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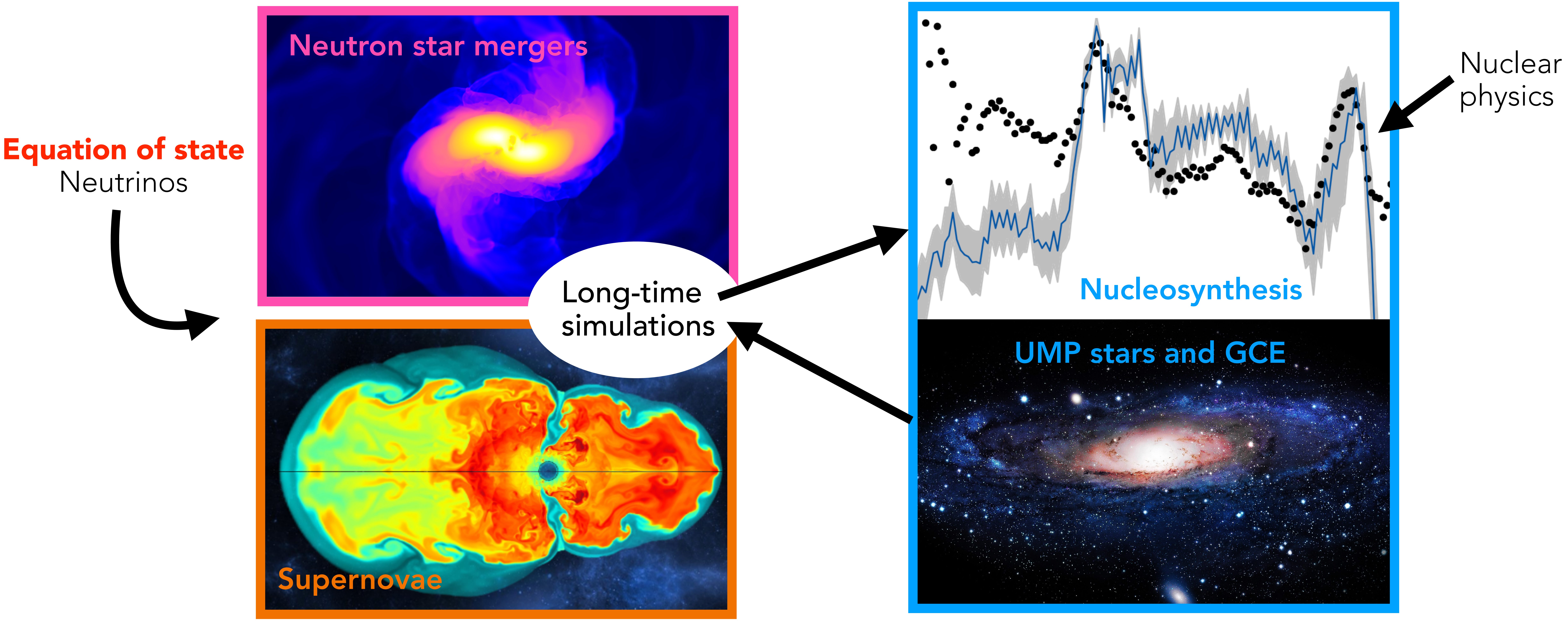
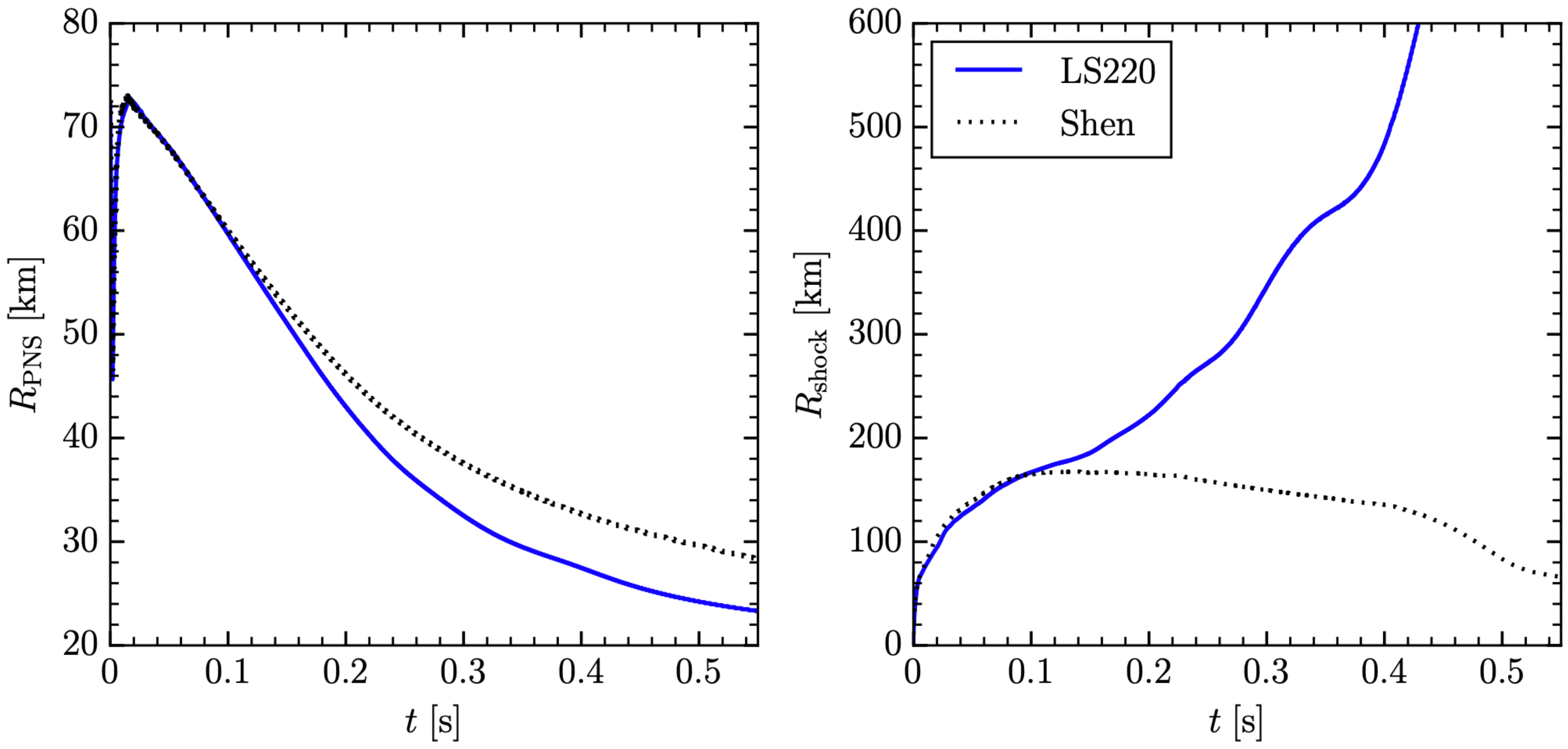


Equation of state in core-collapse supernovae

First systematic study of nuclear matter properties

1D simulations, FLASH + M1 + increased neutrino heating



Yasin et al., PRL (2020)

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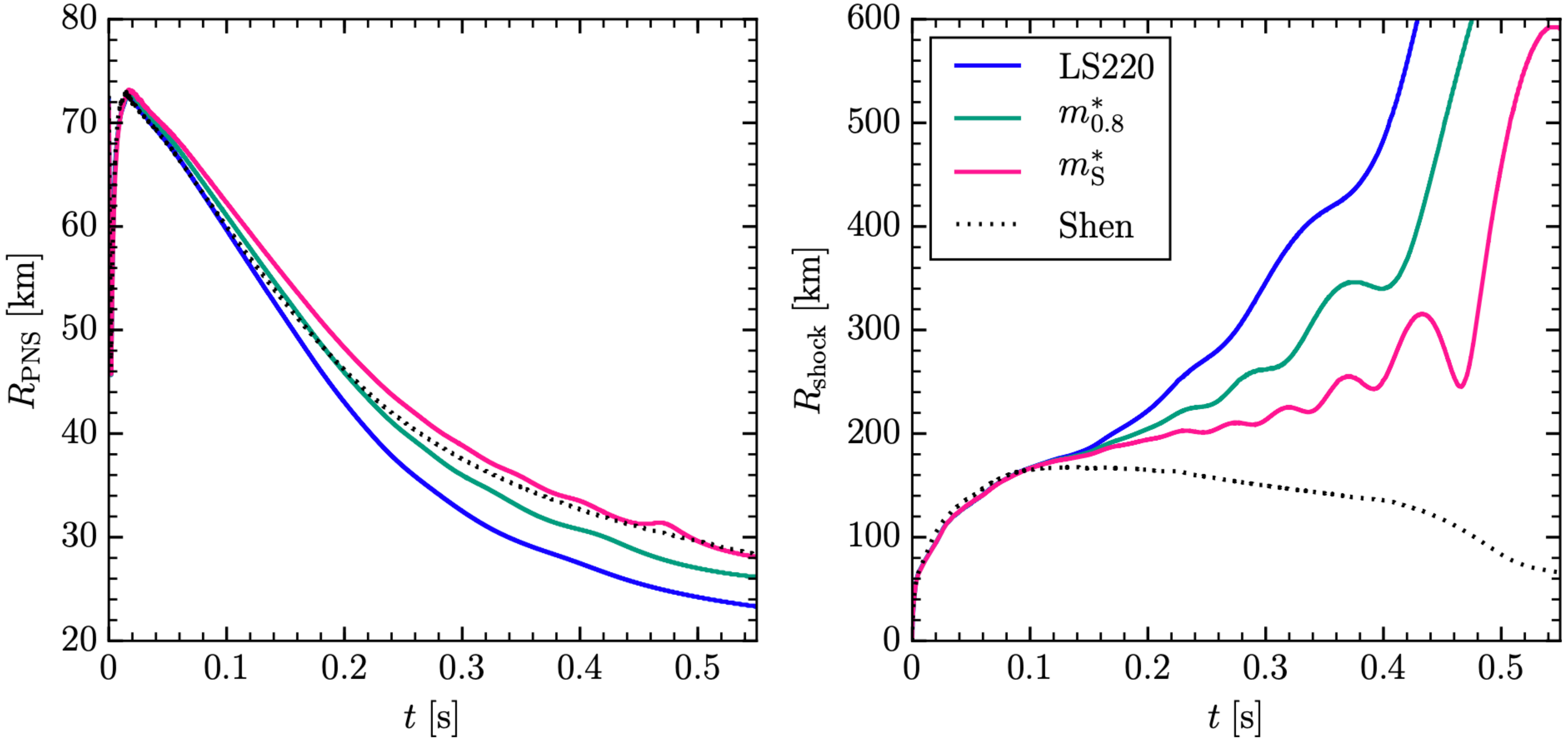

Effective mass:
PNS contraction

Yasin et al., PRL (2020)

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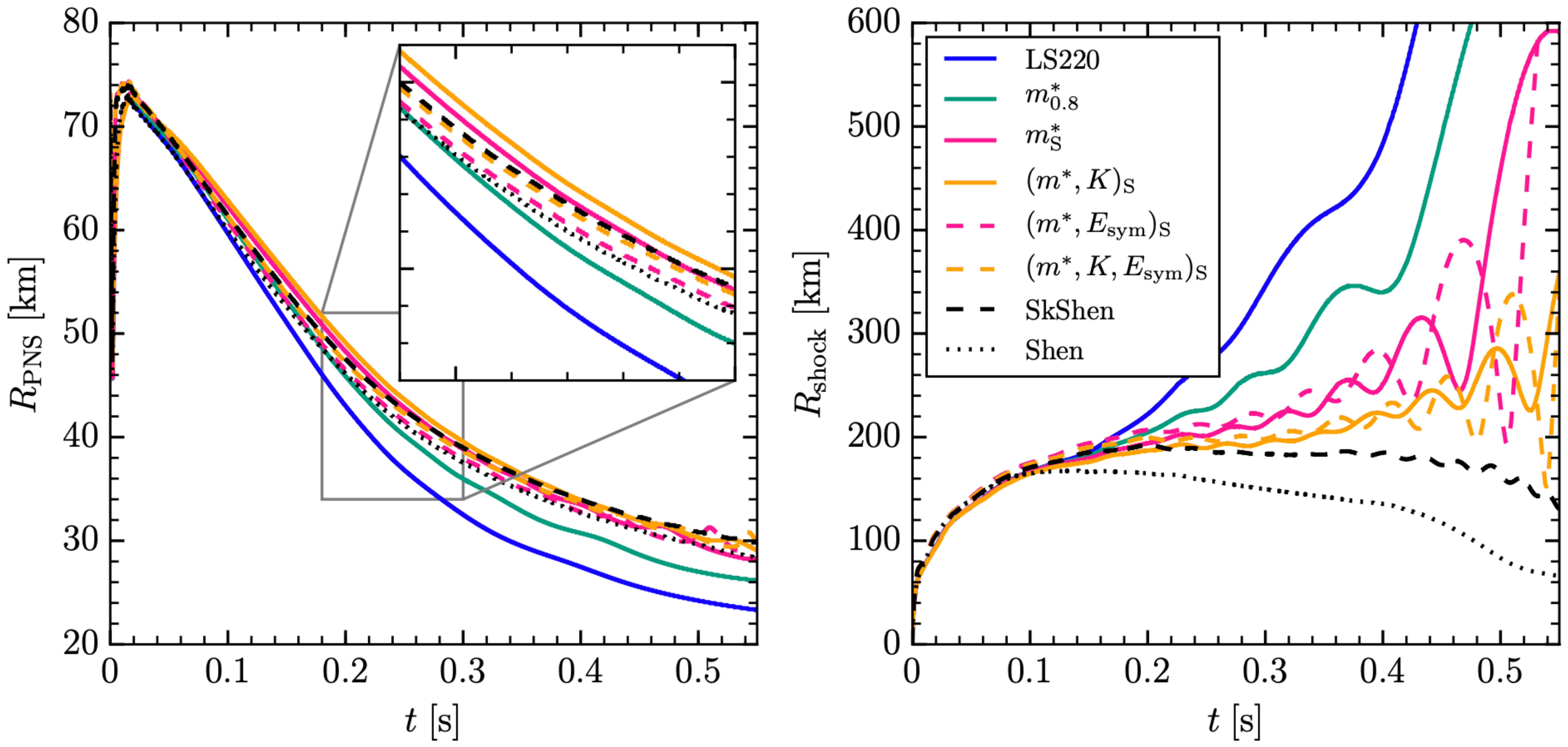

Effective mass:

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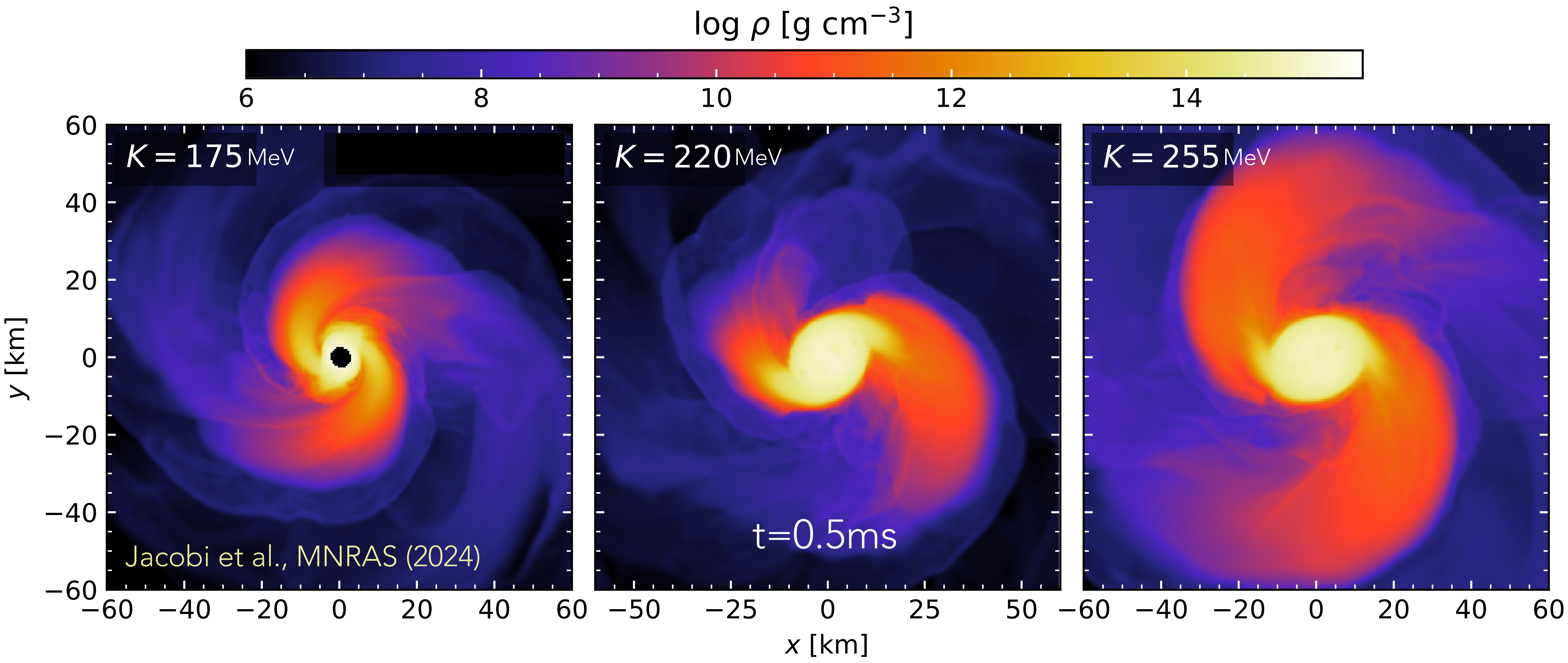
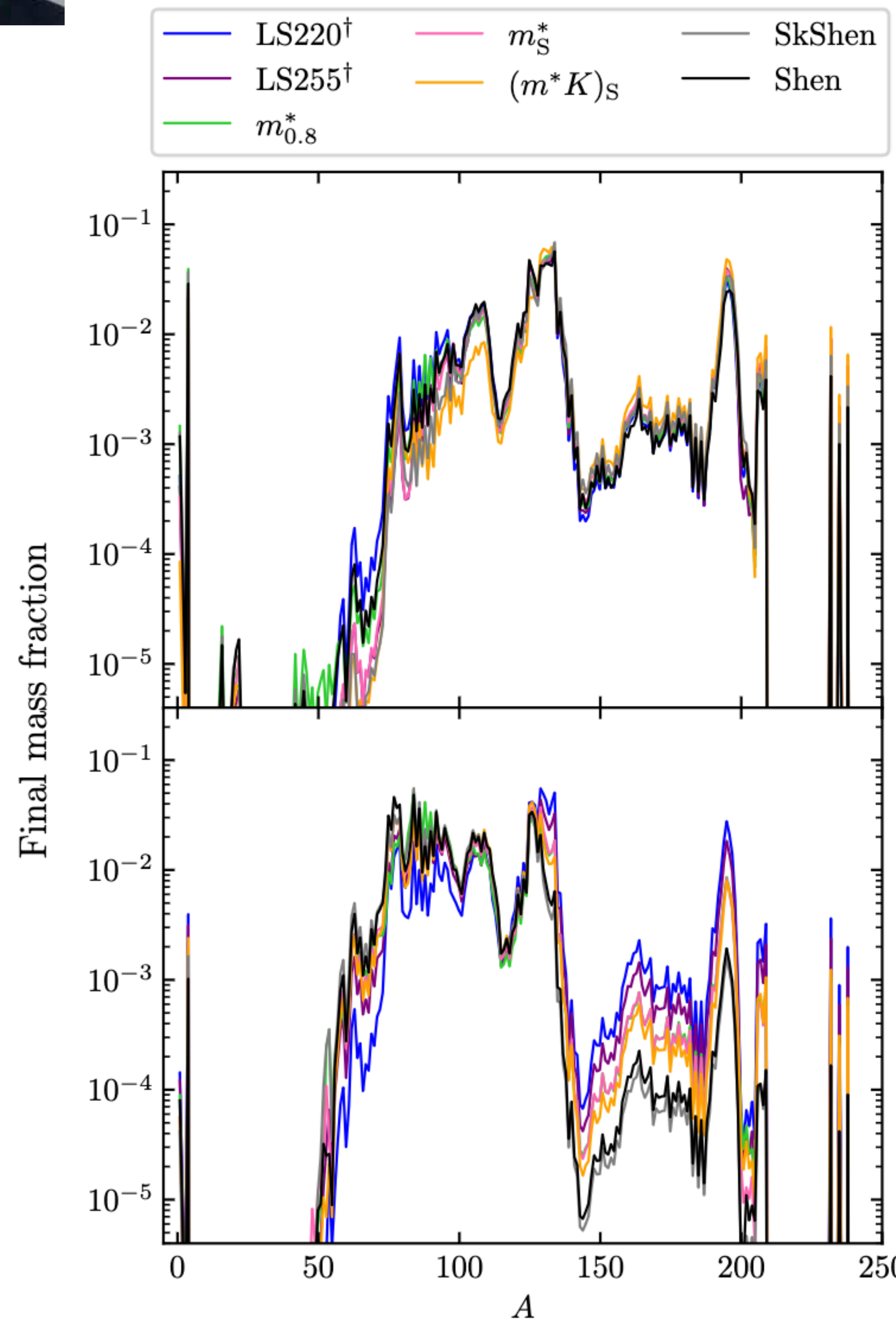
Yasin et al., PRL (2020)

Equation of state in neutron star mergers

Systematic variations of key nuclear matter properties

following Bovard et al. PRC, 2017



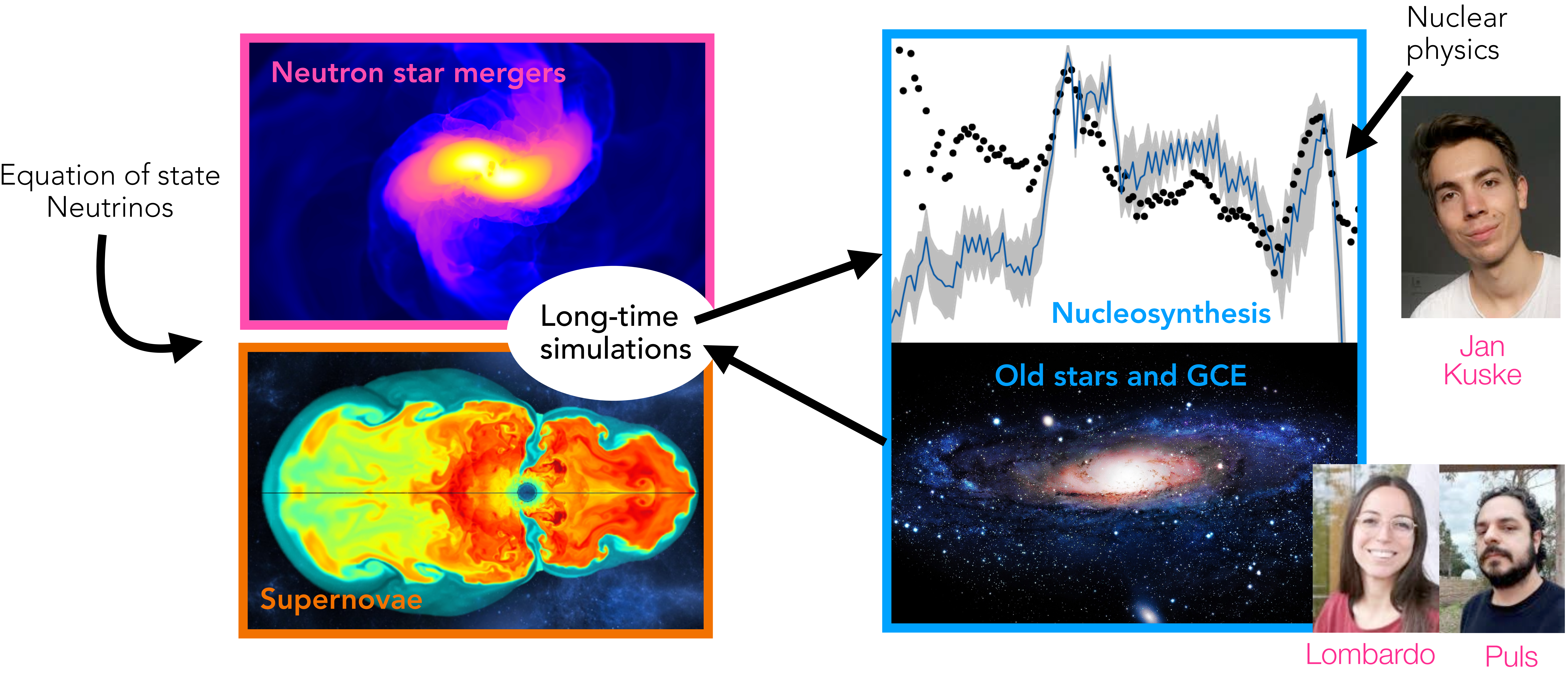


Impact on:

dynamics, gravitational waves, mass ejected (Jacobi et al., MNRAS 2024)

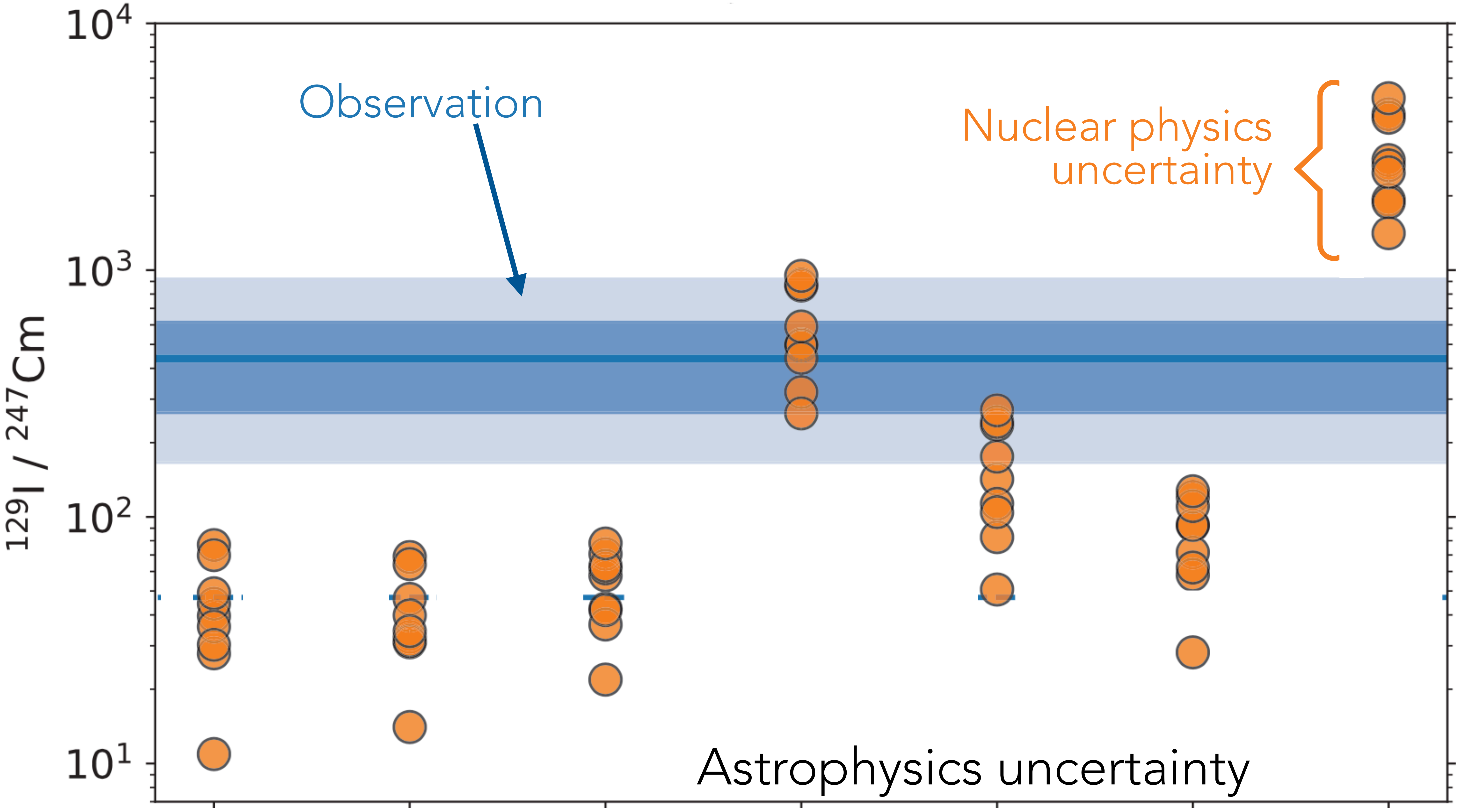
nucleosynthesis and kilonova (Ricigliano et al., MNRAS 2024)

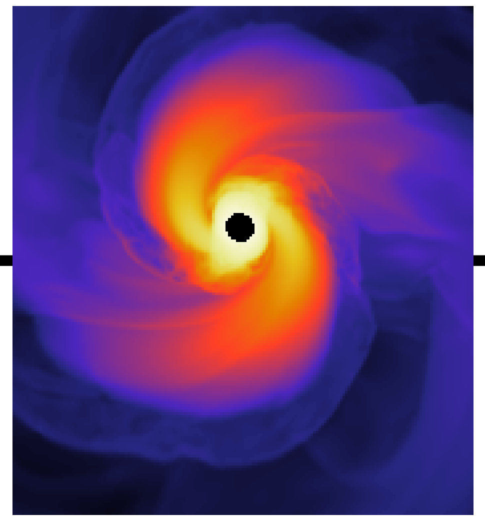
R-process: from simulations to observations

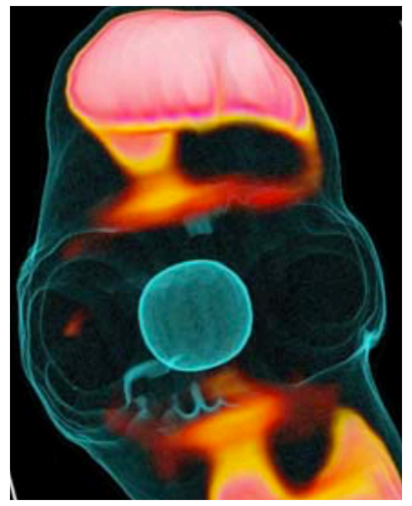


R-process: from simulations to observations

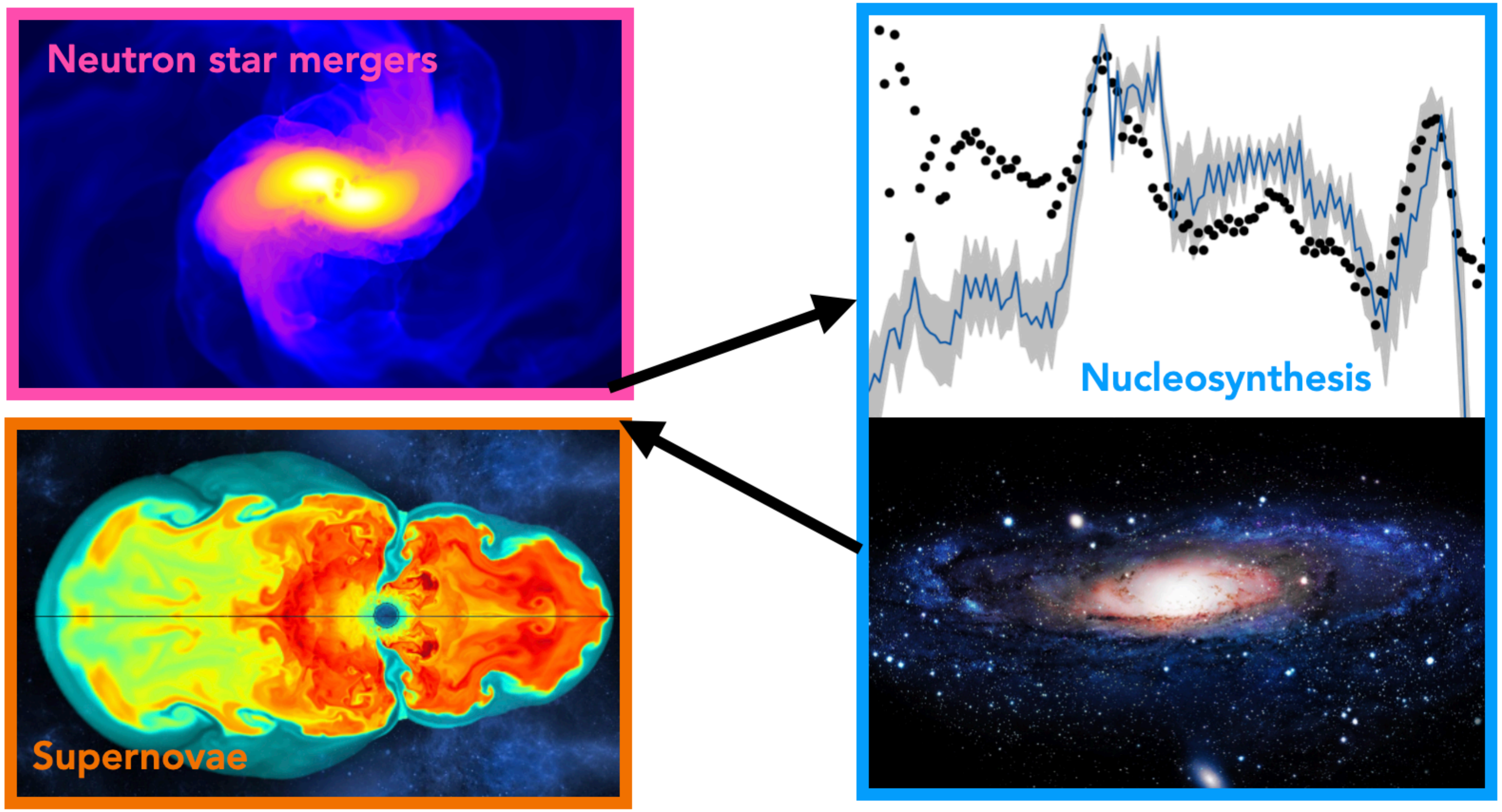
Côtè et al. Science (2021)







Mergers and supernovae as cosmic laboratories establish the origin and history of heavy elements in the universe



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DFG Deutsche Forschungsgemeinschaft